

## THE SPLASH SURVEY: SPECTROSCOPY OF 15 M31 DWARF SPHEROIDAL SATELLITE GALAXIES\*

ERIK J. TOLLERUD<sup>1</sup>, RACHAEL L. BEATON<sup>2,11</sup>, MARLA C. GEHA<sup>3</sup>, JAMES S. BULLOCK<sup>1</sup>, PURAGRA GUHATHAKURTA<sup>4</sup>,  
JASON S. KALIRAI<sup>5,6</sup>, STEVE R. MAJEWSKI<sup>2,11</sup>, EVAN N. KIRBY<sup>7,12</sup>, KAROLINE M. GILBERT<sup>8,12</sup>, BASILIO YNIGUEZ<sup>1</sup>,  
RICHARD J. PATTERSON<sup>2,11</sup>, JAMES C. OSTHEIMER<sup>2,11</sup>, JEFF COOKE<sup>9</sup>, CLAIRE E. DORMAN<sup>4</sup>,  
ABRAR CHOUDHURY<sup>10</sup>, AND MICHAEL C. COOPER<sup>1,12</sup>

<sup>1</sup> Center for Cosmology, Department of Physics and Astronomy, The University of California, Irvine, Irvine, CA, 92697, USA;  
[etolleru@uci.edu](mailto:etolleru@uci.edu), [bullock@uci.edu](mailto:bullock@uci.edu), [byniguez@uci.edu](mailto:byniguez@uci.edu), [m.cooper@uci.edu](mailto:m.cooper@uci.edu)

<sup>2</sup> Department of Astronomy, University of Virginia, P.O. Box 3818, Charlottesville, VA 22903, USA;  
[rlb9n@virginia.edu](mailto:rlb9n@virginia.edu), [srm4n@mail.astro.virginia.edu](mailto:srm4n@mail.astro.virginia.edu), [ricky@virginia.edu](mailto:ricky@virginia.edu), [james.ostheimer@gmail.com](mailto:james.ostheimer@gmail.com)

<sup>3</sup> Astronomy Department, Yale University, New Haven, CT 06510, USA; [marla.geha@yale.edu](mailto:marla.geha@yale.edu)

<sup>4</sup> University of California Observatories/Lick Observatory, University of California at Santa Cruz, Santa Cruz, CA 95064, USA;  
[raja@ucolick.org](mailto:raja@ucolick.org), [cdorman@ucolick.org](mailto:cdorman@ucolick.org)

<sup>5</sup> Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA; [jkalirai@stsci.edu](mailto:jkalirai@stsci.edu)

<sup>6</sup> Center for Astrophysical Sciences, Johns Hopkins University, Baltimore, MD, 21218

<sup>7</sup> Department of Astronomy, California Institute of Technology, 1200 East California Boulevard, MC 249-17,  
Pasadena, CA 91125, USA; [enk@astro.caltech.edu](mailto:enk@astro.caltech.edu)

<sup>8</sup> Department of Astronomy, University of Washington, Box 351580, Seattle, WA 98195, USA; [kmgilber@u.washington.edu](mailto:kmgilber@u.washington.edu)

<sup>9</sup> Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Mail H39, PO Box 218, Hawthorn,  
VIC 3122, Australia; [jcooke@astro.swin.edu.au](mailto:jcooke@astro.swin.edu.au)

<sup>10</sup> Bellarmine College Preparatory, 960 West Hedding Street, San Jose, CA 95126, USA

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### ABSTRACT

We present a resolved star spectroscopic survey of 15 dwarf spheroidal (dSph) satellites of the Andromeda galaxy (M31). We filter foreground contamination from Milky Way (MW) stars, noting that MW substructure is evident in this contaminant sample. We also filter M31 halo field giant stars and identify the remainder as probable dSph members. We then use these members to determine the kinematical properties of the dSphs. For the first time, we confirm that And XVIII, XXI, and XXII show kinematics consistent with bound, dark-matter-dominated galaxies. From the velocity dispersions for the full sample of dSphs we determine masses, which we combine with the size and luminosity of the galaxies to produce mass–size–luminosity scaling relations. With these scalings we determine that the M31 dSphs are fully consistent with the MW dSphs, suggesting that the well-studied MW satellite population provides a fair sample for broader conclusions. We also estimate dark matter halo masses of the satellites and find that there is no sign that the luminosity of these galaxies depends on their dark halo mass, a result consistent with what is seen for MW dwarfs. Two of the M31 dSphs (And XV, XVI) have estimated maximum circular velocities smaller than  $12 \text{ km s}^{-1}$  (to  $1\sigma$ ), which likely places them within the lowest-mass dark matter halos known to host stars (along with Boötes I of the MW). Finally, we use the systemic velocities of the M31 satellites to estimate the mass of the M31 halo, obtaining a virial mass consistent with previous results.

**Key words:** dark matter – galaxies: dwarf – galaxies: fundamental parameters – galaxies: individual (And I, And III, And V, And VII, And IX, And X, And XI, And XII, And XIII, And XIV, And XV, And XVI, And XVIII, And XXI, And XXII) – galaxies: kinematics and dynamics – Local Group

*Online-only material:* color figures, machine-readable table

### 1. INTRODUCTION

Dwarf spheroidal (dSph) galaxies are among the most extreme objects in the pantheon of galaxies. Their low luminosities ( $10^3 < L/L_\odot < 10^8$ ), lack of significant gas (Grcevich & Putman 2009), and low numbers compared to  $\Lambda$ CDM expectations (Klypin et al. 1999; Moore et al. 1999) are all puzzles that remain to be solved. The difficulty in understanding the count of dSph galaxies around the Milky Way (MW) and M31 is known as the missing satellite problem, an issue that has prompted a

flurry of activity modeling these galaxies (recently, Kravtsov 2010; Bullock 2010; Kazantzidis et al. 2011; Font et al. 2011, and references therein). Most models rely heavily on feedback scenarios that are tied directly to the masses of the dark matter halos that presumably host dSph galaxies. In this sense, mass determinations for dwarfs are among the most important diagnostic measurements for testing theoretical predictions at the frontier of galaxy formation.

While the brightest dSphs can be detected at the distance of nearby clusters (e.g., Hilker et al. 2003; Durrell et al. 2007), faint, diffuse dSphs can only be detected via resolved star counts, which limits detection to the Local Group (LG). Kinematics of these galaxies thus require resolved star spectroscopy at extragalactic distances. Thus, despite the motivations outlined above, detailed study of a large population of dSphs has been limited to the MW satellites (e.g., Mateo 1998; Simon & Geha 2007; Walker et al. 2009; Simon et al. 2011).

Studies of the MW dSph population have resulted in puzzles that have provided interesting challenges to  $\Lambda$ CDM and galaxy

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<sup>11</sup> Visiting Astronomer, Kitt Peak National Observatory, National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

<sup>12</sup> Hubble Fellow.



**Table 1**  
Observed DEIMOS Slitmasks

Target	Mask Name	MJD	UTC Date	Total Exposure Time (s)	No. of Slits	No. of Successful Velocities <sup>a</sup>
And I	d1_1	53679.31	2005 Nov 5	4055	150	71
And I	d1_2	53994.62	2006 Sep 16	3600	145	88
And III	d3_1	53621.34	2005 Sep 8	3600	119	88
And III	d3_2	53621.39	2005 Sep 8	3600	117	83
And III	d3_3	55066.56	2009 Aug 23	3600	123	30
And V	d5_1	54739.45	2008 Sep 30	3000	101	52
And V	d5_2	54739.55	2008 Sep 30	2250	105	64
And V	d5_3	54740.54	2008 Oct 1	2250	75	37
And VII	d7_1	54682.43	2008 Aug 4	3000	152	110
And VII	d7_2	54682.46	2008 Aug 4	1800	147	74
And IX	d9_1	55069.56	2009 Aug 26	2700	115	74
And IX	d9_2	55069.60	2009 Aug 26	2400	101	58
And X	d10_1	53618.39	2005 Sep 5	3600	82	51
And X	d10_2	53618.44	2005 Sep 5	3600	93	54
And XI	d11	55450.60	2010 Sep 11	3600	71	22
And XII	d12	55451.00	2010 Sep 12	3600	62	22
And XIII	d13_1	54739.37	2008 Sep 30	6000	66	18
And XIII	d13_2	54740.48	2008 Oct 1	4800	65	15
And XIII	d13_3	55450.31	2010 Sep 11	3600	107	18
And XIII	d13_4	55450.36	2010 Sep 11	5400	114	12
And XIII	d13_5	55450.44	2010 Sep 11	5400	115	11
And XIV	A170_1	54059.20	2006 Nov 20	3600	93	39
And XIV	A170_2	54060.20	2006 Nov 21	3600	88	36
And XIV	d14_3	55065.55	2009 Aug 22	3600	99	33
And XV	d15_1	55068.36	2009 Aug 25	3600	105	82
And XV	d15_2	55068.43	2009 Aug 25	4800	120	60
And XVI	d16_1	55068.50	2009 Aug 25	4800	92	29
And XVI	d16_2	55068.56	2009 Aug 25	3600	77	16
And XVIII	d18_1	55069.40	2009 Aug 26	10800	76	49
And XXI	d21_1	55710.56	2011 May 29	6905	96	12
And XXII	d22_1	55450.55	2010 Sep 11	3600	140	10
And XXII	d22_2	55451.56	2010 Sep 12	4800	76	7
And XXII	d22_3	55476.59	2010 Oct 7	3955	125	6

**Notes.** <sup>a</sup> While for some masks there is an apparently low percentage of successful velocities, this is due to higher fractions of filler targets in the sparser fields that tend to have lower odds of being RGB stars. This is discussed in detail in Section 3.1.

This paper is organized as follows. In Section 2, we describe the observations performed for this data set. In Section 3, we describe the reduction and membership analysis performed homogeneously across the spectroscopic data set, as well as our method for estimating total velocity dispersions for each satellite in the sample. In Section 4, we present the results of our full spectroscopic sample and consider each satellite in turn, describing the results of our analysis and unique aspects of each galaxy. In Section 5, we consider the scaling relations of M31 dSphs and compare them to the MW. In Section 6, we use the galaxies'  $v_{\text{sys}}$  to estimate the mass of M31. Finally, we present our conclusions in Section 7.

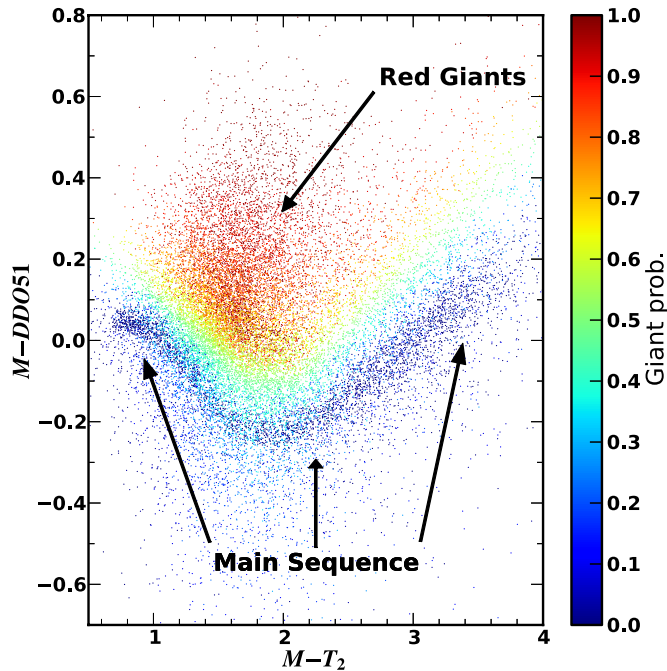
## 2. OBSERVATIONS

We provide an overview of the M31 satellite system in Figure 1. M31 is represented as the (orange) ellipse near the center, while the other symbols are likely M31 satellites. The 15 satellites presented here with Keck/DEIMOS spectroscopy of resolved stars are shown as filled (blue) triangles, while the remainder of the dSphs are shown as outlined (red) triangles (excepting the newly discovered And XXVIII, which is much farther out than the other dSphs). Also shown are M33 (cyan ellipse at the lower left) and the dE satellites (black squares). From this census it is apparent that our sample includes over half

of the dSphs in the M31 system, with the spatial unevenness only due to the very recent discovery of some of the dSphs. Also shown are lines of Galactic latitude, indicating significant variation in the MW foreground across M31's environs.

We present 33 DEIMOS slitmasks covering 15 of M31's dSphs, with mask details provided in Table 1. Eight of these masks have been presented elsewhere: the And I and III masks were previously presented in Kalirai et al. (2010), the And X data set was originally published by Kalirai et al. (2009), and two masks of And XIV were described by Majewski et al. (2007). There is an additional series of SPLASH slitmasks for And II that will be described in a forthcoming paper (N. Ho et al. 2012, in preparation). All masks were reduced and analyzed homogeneously using the procedure described in the following sections, including those noted above that have been previously published.

The details of the imaging used for target selection will be described in detail in the companion paper (R. L. Beaton et al. 2012, in preparation). Here we summarize the particulars relevant for spectroscopic target selection. Our imaging is primarily in the Washington system (specifically, the  $M$  and  $T_2$  filters), which we convert for some target selection purposes to Johnson–Cousins  $V$ ,  $I$  using the relations of Majewski et al. (2000). We obtained additional imaging of the same fields with the DDO51 intermediate-band filter (Majewski et al. 2000).



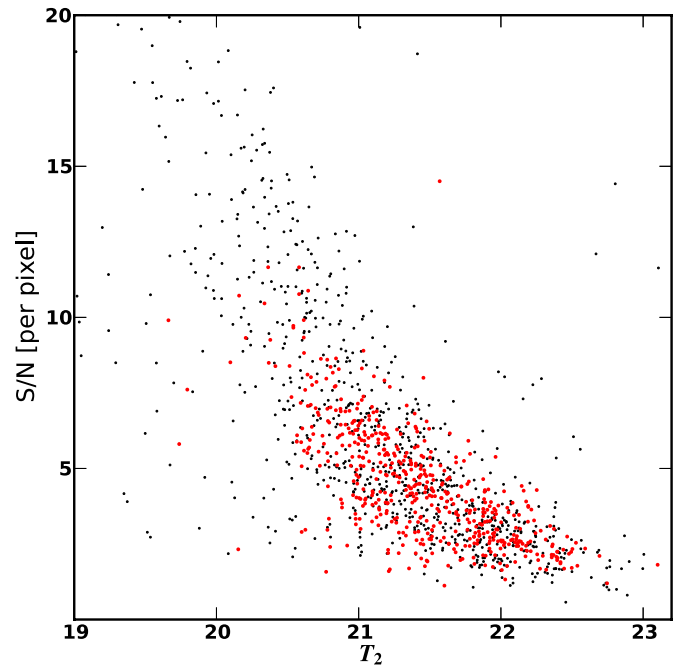
**Figure 2.**  $M - T_2$ ,  $M - \text{DDO51}$  color-color diagram used for pre-selection of likely giant stars, for a representative field near And I. Points are colored by red giant probability determined by the method of Ostheimer (2003), which assigns high probabilities to stars near the indicated locus for stars with low surface gravity, while avoiding probable dwarf stars in the “swoosh” feature. Stars with high probability are selected for spectroscopic follow-up.

(A color version of this figure is available in the online journal.)

This filter is centered near the surface-gravity-dependent MgB and MgH stellar absorption features, allowing it to discriminate between M31 giant stars and MW foreground dwarf stars. Thus, selecting targets in, e.g., the  $M - \text{DDO51}$  versus  $M - T_2$  color-color diagram allows for far more efficient selection of spectroscopic targets that have high probabilities of being giants (Kalirai et al. 2006; Guhathakurta et al. 2006; Gilbert et al. 2006).

We selected stars for spectroscopic follow-up from catalogs generated from this imaging. Our candidates for each mask were selected by assigning priorities based on how far a given star was from a fiducial isochrone in the  $M - T_2$ ,  $T_2$  color-magnitude diagram (CMD). This is similar to the membership selection described in Section 3.2 and illustrated in the plots of Section 4, but with looser restrictions. We also assigned priorities based on how close a star was to the giant star locus in the  $M - \text{DDO51}$  versus  $M - T_2$  color-color diagram (Ostheimer 2003). We demonstrate this part of the selection in Figure 2, showing stars for a representative field and indicating the loci of red giant stars (the targets) and main-sequence stars (the contaminants), with points colored according to an estimate of the probability of the star being a giant. Because most M31 dSphs are smaller than the footprint of a DEIMOS slitmask ( $\approx 16' \times 5'$ ), in most cases we aligned the mask such that the dSph was centered to one side of the slitmask. This allows for measurement at larger radii of M31 halo structure within which a satellite might be embedded (e.g., And I as described by Kalirai et al. 2010).

Note that the above description is valid for all masks except those targeted at And X, And XV, And XVI, And XVIII, and And XXII. For And X, Sloan Digital Sky Survey (SDSS) imaging was used for target selection (Adelman-McCarthy et al. 2006), for And XV and XVI selection was made using



**Figure 3.** Per-pixel signal-to-noise ratio vs.  $T_2$ -band magnitude for stars in this survey with successful spectra. Larger (red) circles are stars identified as likely dSph members following the prescription described in Section 3.2, while smaller (black) points are non-member stars.

(A color version of this figure is available in the online journal.)

Canada–France–Hawaii Telescope archival imaging, while for the last two dSphs we made use of  $B$ - and  $V$ -band imaging from the Large Binocular Telescope (LBT; R. L. Beaton et al. 2012, in preparation). For these, distance from a fiducial isochrone in a CMD was used without DDO51 pre-selection.

The spectroscopic setup for the DEIMOS observations used the 1200 line  $\text{mm}^{-1}$  grating with a central wavelength of 7800 Å. This provides spectral coverage over a range of 6400–9100 Å (for objects centered in the mask along the dispersion direction), with an FWHM resolution of  $\approx 1.3$  Å. This provides coverage of the calcium triplet (CaT) stellar absorption feature near 8500 Å and (depending on the slit location) H $\alpha$  to facilitate identification and accurate radial velocity measurements. A typical total integration time of 3600 s per mask provides a mean signal-to-noise ratio (S/N) per pixel of  $\sim 7$  for our entire sample. S/N depends strongly on the target star’s magnitude, so we plot in Figure 3 the variation of S/N with magnitude.

### 3. SPECTROSCOPIC DATA ANALYSIS

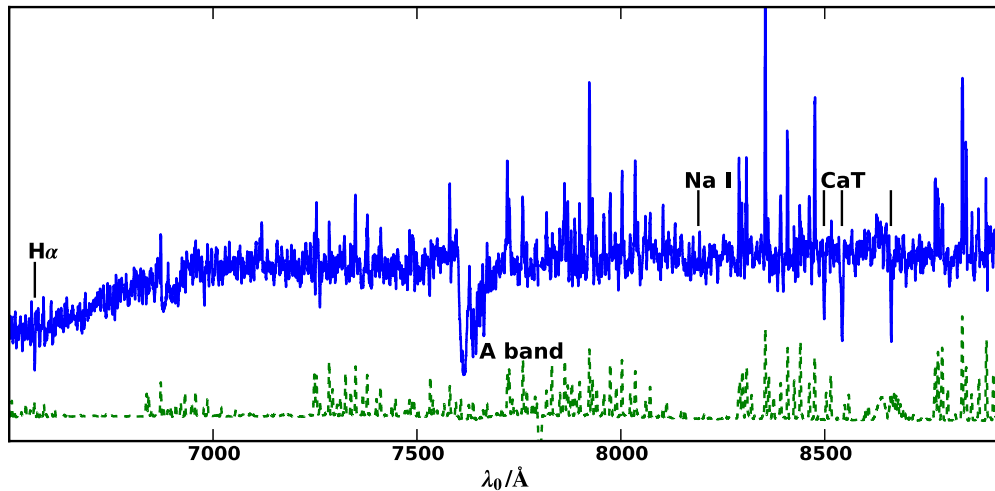
Here we describe the procedure applied homogeneously to each slitmask (Table 1) to extract radial velocities from the spectroscopic data and to determine each satellite’s kinematics.

#### 3.1. Keck/DEIMOS Data Reductions

We make use of the spec2d DEIMOS reduction pipeline<sup>13</sup> and modified versions of the spec1d and zspec analysis codes developed for the DEEP2 survey (Davis et al. 2003; Cooper et al. 2012; Newman et al. 2012). To summarize, the pipeline first rectifies each slit spectrum along the spatial direction and corrects for the slit function and fringing using quartz lamp flats. It then determines a two-dimensional wavelength solution from

<sup>13</sup> <http://deep.berkeley.edu/spec2d/>





**Figure 4.** Example spectrum of a star with a high probability of dSph membership. The solid (blue) line is the spectrum (smoothed with a 4 pixel boxcar filter), while dashed (green) shows the error spectrum, peaking at strong sky emission lines. For this star,  $I = 20.7$ , and un-smoothed S/N =  $6 \text{ pixel}^{-1}$ .

(A color version of this figure is available in the online journal.)

NeArKrXe arc lamp exposures taken for each mask. Science exposures are then combined with inverse variance weighting including cosmic-ray rejection. Finally, one-dimensional (1D) spectra are extracted using Horne (1986) optimal extraction.

The above procedure thus yields 1D spectra for each slit, an example of which is given in Figure 4. These spectra are then further analyzed to determine heliocentric radial velocities, following the same general procedure described by Simon & Geha (2007; see also Guhathakurta et al. 2006; Kalirai et al. 2010). The spectra are first re-binned into equally spaced logarithmic bins of  $15 \text{ km s}^{-1}$ . These spectra are then cross-correlated with high-S/N templates taken with DEIMOS using an identical spectroscopic setup. These templates include dwarf, subgiant, giant, and asymptotic giant branch stars. The largest subset is the giant templates, which range from spectral type F8 to M8, although most of our likely members best match early-K and late-G templates.

We then manually examine each spectrum and remove contaminant spectra that are clearly galaxies, as well as those for which no absorption features are visible, i.e., low-S/N spectra. This process results in very different success rates from mask to mask because masks with a low density of good M31 red giant branch (RGB) candidates will include a large number of “filler” targets. These targets tend to be fainter by necessity, meaning the photometric errors are larger, and the fraction of background galaxies that overlap with the M31 RGB is much higher, increasing the contamination rate. Furthermore, even true RGB targets are fainter, and thus identification of absorption lines becomes more difficult, further lowering the success rate. Together, these effects are the main reasons for the high variability in mask-to-mask success rates (i.e., the ratio of the last two columns in Table 1).

We also correct for mis-centering of stars in slits by matching to the A-band telluric feature as described by Simon & Geha (2007) and Sohn et al. (2007). This correction is crucial, as mis-centering of stars in slits can result in velocity scatter comparable to that of the internal velocity dispersion of the dSphs for typical astrometric errors (see the discussion of Kalirai et al. 2010, Section 2.3). Finally, we correct the observed radial velocities to the heliocentric frame.

Understanding the errors on the radial velocity computed above is crucial to determining the internal velocity dispersion

of the dSphs, and hence we determine the errors on each star by Monte Carlo simulation. For each spectrum that passes the inspection procedure described above, we generate 1000 realizations of the 1D spectrum, assuming independent, Poisson errors on each pixel based on the observed variance. For each of these realizations we compute the telluric correction and the cross-correlation radial velocity and visually examine the histogram of radial velocities for these simulations. We reject any spectra that show multi-modal character in radial velocity (e.g., a secondary cross-correlation peak that is of comparable magnitude to the primary peak). We then fit a Gaussian to the primary peak and use the best-fit parameters as the radial velocity  $v_i$  and random error  $\sigma_r$  for each spectrum. Repeat measurements of stars for this spectroscopic setup are described in detail by Simon & Geha (2007) and Kalirai et al. (2010, Section 2.4), and our sample includes the same repeat measurements as the latter. Both sets of repeat measurements showed an additional systematic error component of  $2.2 \text{ km s}^{-1}$ . Thus, we adopt  $\sigma_t \equiv \sqrt{\sigma_r^2 + (2.2 \text{ km s}^{-1})^2}$  as our final radial velocity error estimate for each spectrum.

We also measure equivalent widths (EWs) for the Na I  $\lambda 8190$  doublet ( $\lambda\lambda 8183, 8195$ ), a surface-gravity-dependent feature useful for dwarf/giant discrimination. We adopt a blue pseudocontinuum from 8155–8175 Å, red 8203–8230 Å, and a line bandpass of 8178–8200.5 Å (Schiavon et al. 1997). For these measurements, we determine errors using the same Monte Carlo realizations and Gaussian fitting procedure described above for radial velocities.

This procedure thus provides a sample of stellar spectra with heliocentric radial velocities (and corresponding error estimates), as well as indices for spectral features that are typically detectable for red giants at the distance of M31. We match these spectra to the corresponding photometric measurements to create a catalog of candidate M31 dSph stars. This provides the input to our membership selection scheme, described below.

### 3.2. Membership

After determining the radial velocity for each star in our sample, we next determine whether it is a member of a dSph galaxy or a contaminant. There are two distinct populations

of stellar contaminants: MW red dwarf stars (foreground) and M31 red giants (halo field stars). We consider each of these in turn and then describe our method for assigning membership probabilities to individual stars.

Foreground contamination from MW stars is typically worse than the M31 component, particularly for dwarfs that lie closer to the Galactic plane than M31. Gilbert et al. (2006) describe a step-by-step approach for empirically determining the probability that a star is an MW contaminant or an M31 RGB star. While such an approach is valuable for searching for low-density M31 halo field stars, the dSphs we consider here have much higher surface densities in their centers, and thus these diagnostics are less crucial near the centers of the satellite. Furthermore, given that the dSphs studied here exhibit low-dispersion kinematical “cold spikes” at velocities closer to M31’s systemic velocity than that of the MW, even a wide velocity window that extends into the tails of the satellite velocity distribution will overlap negligibly with the MW stars. This is borne out in the velocity histograms shown in Section 4, as the distributions from the Robin et al. (2003) model of the MW do not overlap with the velocity windows we adopt.

Halo field and spheroid stars from M31 also provide a significant source of contamination. While these are generally lower surface density than MW foreground, M31 contamination is more problematic for selecting dSphs due to two effects. First, the DDO51 pre-selection technique described above is only effective at eliminating foreground dwarf stars—it cannot distinguish M31 dSph red giants from M31 halo field red giants. Second, while the velocity distribution of the M31 halo is much broader than that of the dSphs, it overlaps with the dSph distributions for many of the satellites. Hence, there is no definitive way to exclude M31 halo field stars from the dSphs with radial velocity measurements. Indeed, given the hypothesis that the outer reaches of galactic halos or spheroids are formed from dissolved satellites (e.g., Bullock & Johnston 2005; Guhathakurta et al. 2006), it is unsurprising that the halo’s stellar properties overlap with those of the satellites.

Fortunately, the surface densities of M31 halo field stars in the environs of these satellites are quite low relative to the dSphs. M31’s spheroid (non-disk) component can be decomposed into two sub-components: an “inner spheroid” that extends to  $\sim 30$  kpc and features a de Vaucouleurs profile, and a power-law “halo” component with a slope of  $\sim -2.3$  extending out to at least 160 kpc (Guhathakurta et al. 2005). All of our dSph fields, aside from And I and IX (see Sections 4.1 and 4.5), lie firmly in the halo regime, and the distances imply at most a few halo field stars per DEIMOS slitmask (Guhathakurta et al. 2005, 2006; Gilbert et al. 2007). Hence, the probability of finding such a star near the center of a dSph with a similar velocity to that of the member stars is low. Additionally, some of the dSphs have systemic velocities well separated from the M31 spheroid—the M31 halo distribution is centered at  $\sim -300$  km s $^{-1}$  with a dispersion of  $\sim 100$  km s $^{-1}$  (Guhathakurta et al. 2005, 2006; Chapman et al. 2006; Gilbert et al. 2007). Furthermore, the metallicity distribution functions of the M31 halo and the dSphs are partially disjoint. M31 dSphs examined here tend to have metallicities in the range  $[\text{Fe}/\text{H}] \sim -1.4$  to  $-2.0$  (Martin et al. 2006, 2009; McConnachie et al. 2008; Kalirai et al. 2010), while the M31 spheroid is more concentrated in the range  $[\text{Fe}/\text{H}] \sim -1.5$  to  $-0.1$  (Guhathakurta et al. 2005, 2006; Chapman et al. 2006). While there is overlap, particularly in the outer halo where M31 metallicities are lower, the distributions are reasonably separated in the mean.

Thus, metallicity-based criteria have hope of approximately separating M31 halo structure from dSph stars in the outskirts of the satellites. While we do not make use of spectroscopic metallicities due to the lower S/N for some of our spectra, our selection depends on this metallicity distinction implicitly through CMD selection.

With these two sources of contamination in mind, we adopt a method for determining membership probabilities that depends only on distance from the center of the target satellite and location in the  $(T_2, M - T_2)$  CMD. Explicitly, the membership probability for each star in the survey is calculated as

$$P_{\text{memb}} = \exp \left( -\frac{\Delta\alpha^2 + \Delta\delta^2}{2\eta R_{\text{eff}}^2} - \frac{\Delta(M - T_2)^2}{2\sigma_c} - \frac{\Delta T^2}{2\sigma_m} \right) \times \left( \frac{\arctan[-10(\Sigma_{\text{Na}} - 2)]}{\pi} + \frac{1}{2} \right), \quad (1)$$

where  $\Delta\alpha$  and  $\Delta\delta$  are the circular distance from the star to the center of the dSph (in R.A. and decl., respectively),  $\Delta(M - T_2)$  and  $\Delta T$  are distances from fiducial isochrones in the  $(T_2, M - T_2)$  CMD,  $R_{\text{eff}}$  is the dSph’s half-light radius,  $\Sigma_{\text{Na}}$  is the Na I EW, and  $\eta$ ,  $\sigma_c$ , and  $\sigma_m$  are free parameters.

The first term of this expression simply assigns a circular Gaussian spatial distribution of probabilities, weighting the region inside the half-light radius. While some of the dSphs have elliptical light distributions, here we use a circular distribution to avoid introducing biases for the cases where the ellipticity is not well measured. In most cases, we adopt  $\eta = 1.5$ , except in cases where halo contamination is much different from typical as indicated by the velocity distribution at large distances from the satellite. In these cases we use a slightly different  $\eta$  value to down- or up-weight the membership probability for stars beyond  $R_{\text{eff}}$ . The satellites for which this is done are explicitly mentioned in Section 4.

The latter two terms inside the exponential define Gaussian acceptance regions around the fiducial isochrone. We make use of Girardi et al. (2002) isochrones<sup>14</sup> in the Washington photometric system for a range of metallicities from  $[\text{Fe}/\text{H}] = -2.2$  to  $-1.0$  with 12 Gyr ages, offset by the distance modulus of the satellite. We then correct the  $M$  and  $T_2$  photometry for Galactic extinction/reddening using the Schlegel et al. (1998) dust maps and use photometry of the DDO51-selected stars within each dwarf’s half-light radius to empirically choose the best-fitting isochrone. With these likely members as a guide, we adopt fiducial values of  $\sigma_c = 0.1$  and  $\sigma_m = 0.5$ , although these parameters are adjusted for some dwarfs, as described in Section 4.

Given that the only stars we can reach spectroscopically at the distance of M31 are red giant stars, which form a nearly vertical sequence in the CMD, the range of accepted colors, encoded in  $\sigma_c$ , contains contributions from both intrinsic metallicity scatter and photometric errors. Meanwhile,  $\sigma_m$  contains both distance and photometric uncertainties. Uncertainty in the age of the stellar population induces further degeneracy in metallicity and distance. This motivates the choice of fitting isochrones empirically rather than performing a detailed analysis of the isochrone-inferred metallicities, leaving such discussion to the companion paper (R. L. Beaton et al. 2012, in preparation). We instead use the isochrones primarily to reject M31 halo field stars (which have a disjoint metallicity distribution, as described

<sup>14</sup> <http://stev.oapd.inaf.it/cgi-bin/cmd>

**Table 2**  
SPLASH M31 dSph Stars

Mask <sup>a</sup>	Star ID <sup>b</sup>	R.A. <sup>c</sup>	Decl. <sup>d</sup>	$T_2$ <sup>e</sup>	$M - T_2$ <sup>f</sup>	$v_{\text{helio}}^g$ (km s <sup>-1</sup> )	Na EW <sup>h</sup>	$P_{\text{memb}}^i$
d1_1	1005593	0 <sup>h</sup> 45 <sup>m</sup> 47 <sup>s</sup> .47	+38°6′30″.5	21.725 ± 0.039	1.46 ± 0.06	−400.0 ± 6.3	1.2 ± 1.4	0.141
d1_1	1007403	0 <sup>h</sup> 45 <sup>m</sup> 48 <sup>s</sup> .97	+38°6′05″.0	20.572 ± 0.015	1.82 ± 0.03	−373.6 ± 2.6	0.3 ± 0.5	0.182
d1_1	1005564	0 <sup>h</sup> 45 <sup>m</sup> 44 <sup>s</sup> .71	+38°6′59″.1	20.735 ± 0.018	1.80 ± 0.03	−378.6 ± 2.4	0.5 ± 0.5	0.11
...	...	...	...	...	...	...	...	...

**Notes.** Stars not classified as members but with successful radial velocity measurements can be made available upon request to the authors

<sup>a</sup> Name of the mask on which this star was observed (see Table 1).

<sup>b</sup> ID number of the star (unique within a dSph field).

<sup>c</sup> Right ascension of the star (J2000).

<sup>d</sup> Declination number of the star (J2000).

<sup>e</sup> Washington  $T_2$  magnitude (extinction corrected).

<sup>f</sup> Washington  $M - T_2$  color (extinction corrected).

<sup>g</sup> Heliocentric radial velocity of this star.

<sup>h</sup> Equivalent width of Na I  $\lambda$ 8190 feature.

<sup>i</sup> Membership probability computed following Equation (1).

(This table is available in its entirety in machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

**Table 3**  
Summary of M31 dSph Properties

Name <sup>a</sup>	R.A. <sup>b</sup>	Decl. <sup>c</sup>	$M_V$ <sup>d</sup>	$d_{\text{LOS}}^e$	$r_{\text{M31}}^f$	$N_{\text{memb}}^g$	$v_{\text{sys}}^h$	$\sigma_{\text{LOS}}^i$	Sources <sup>j</sup>
And I	00:45:39.800	+38:02:28.00	−11.8 <sup>+1.0</sup> <sub>−1.0</sub>	744.7 <sup>+24.4</sup> <sub>−23.6</sub>	58.4 <sup>+35.4</sup> <sub>−34.3</sub>	51	−376.3 ± 2.2 <sup>k</sup>	10.2 ± 1.9 <sup>k</sup>	a,b,c
And III	00:35:33.800	+36:29:52.00	−10.2 <sup>+0.3</sup> <sub>−0.3</sub>	748.2 <sup>+24.5</sup> <sub>−23.7</sub>	75.2 <sup>+35.5</sup> <sub>−34.4</sub>	62	−344.3 ± 1.7 <sup>k</sup>	9.3 ± 1.4 <sup>k</sup>	a,c,d
And V	01:10:17.100	+47:37:41.00	−9.6 <sup>+0.3</sup> <sub>−0.3</sub>	820.4 <sup>+15.3</sup> <sub>−15.0</sub>	118.2 <sup>+29.9</sup> <sub>−29.0</sub>	85	−397.3 ± 1.5	10.5 ± 1.1	a,e,f
And VII	23:26:31.700	+50:40:33.00	−13.3 <sup>+0.3</sup> <sub>−0.3</sub>	762.1 <sup>+25.0</sup> <sub>−24.2</sub>	218.3 <sup>+35.8</sup> <sub>−34.7</sub>	136	−307.2 ± 1.3	13.0 ± 1.0	a,c,g
And IX	00:52:53.000	+43:11:45.00	−8.1 <sup>+0.4</sup> <sub>−0.1</sub>	765.6 <sup>+25.1</sup> <sub>−24.3</sub>	40.5 <sup>+35.9</sup> <sub>−34.7</sub>	32	−209.4 ± 2.5	10.9 ± 2.0	a,c,h
And X	01:06:33.700	+44:48:16.00	−7.4 <sup>+0.1</sup> <sub>−0.1</sub>	701.5 <sup>+33.1</sup> <sub>−31.6</sub>	109.4 <sup>+41.9</sup> <sub>−40.2</sub>	27	−164.1 ± 1.7	6.4 ± 1.4	a,i
And XI	00:46:20.000	+33:48:05.00	−6.9 <sup>+0.5</sup> <sub>−0.1</sub>	871.0 <sup>+84.0</sup> <sub>−76.6</sub>	139.1 <sup>+87.9</sup> <sub>−80.6</sub>	2	−461.8 ± 3.7 <sup>l</sup>	...	a,j
And XII	00:47:27.000	+34:22:29.00	−6.4 <sup>+0.1</sup> <sub>−0.5</sub>	831.8 <sup>+47.3</sup> <sub>−44.7</sub>	109.2 <sup>+53.8</sup> <sub>−51.2</sub>	2	−525.3 ± 3.4 <sup>l</sup>	...	a,j,k
And XIII	00:51:51.000	+33:00:16.00	−6.7 <sup>+0.4</sup> <sub>−0.1</sub>	871.0 <sup>+84.0</sup> <sub>−76.6</sub>	150.0 <sup>+87.9</sup> <sub>−80.6</sub>	12	−185.4 ± 2.4	5.8 ± 2.0	a,j
And XIV	00:51:35.000	+29:41:49.00	−8.5 <sup>+0.1</sup> <sub>−0.1</sub>	734.5 <sup>+120.6</sup> <sub>−103.6</sub>	162.3 <sup>+123.3</sup> <sub>−106.5</sub>	48	−480.6 ± 1.2	5.3 ± 1.0	a,l
And XV	01:14:18.700	+38:07:03.00	−9.8 <sup>+0.4</sup> <sub>−0.4</sub>	770.0 <sup>+70.0</sup> <sub>−70.0</sub>	93.6 <sup>+74.6</sup> <sub>−74.3</sub>	29	−323.0 ± 1.4	4.0 ± 1.4	a,m,n
And XVI	00:59:29.800	+32:22:36.00	−9.2 <sup>+0.5</sup> <sub>−0.5</sub>	525.0 <sup>+50.0</sup> <sub>−50.0</sub>	279.4 <sup>+56.2</sup> <sub>−55.8</sub>	7	−367.3 ± 2.8	3.8 ± 2.9	a,m,n
And XVIII	00:02:14.500	+45:05:20.00	−9.7 <sup>+0.1</sup> <sub>−0.1</sub>	1355.2 <sup>+83.6</sup> <sub>−78.8</sub>	590.9 <sup>+87.5</sup> <sub>−82.6</sub>	22	−332.1 ± 2.7	9.7 ± 2.3	a,o
And XXI	23:54:47.700	+42:28:15.00	−9.3 <sup>+0.1</sup> <sub>−0.1</sub>	859.0 <sup>+51.0</sup> <sub>−51.0</sub>	149.2 <sup>+57.1</sup> <sub>−56.7</sub>	6	−361.4 ± 5.8	7.2 ± 5.5	a,p
And XXII	01:27:40.000	+28:05:25.00	−6.2 <sup>+0.1</sup> <sub>−0.1</sub>	794.0 <sup>+239.0</sup> <sub>−239.0</sub>	220.6 <sup>+240.4</sup> <sub>−240.3</sub>	7	−126.8 ± 3.1 <sup>m</sup>	3.54 <sup>+4.16</sup> <sub>−2.49</sub>	a,p

**Notes.**

<sup>a</sup> Name of the dSph.

<sup>b</sup> Right ascension of the dSph.

<sup>c</sup> Declination of the dSph.

<sup>d</sup> V-band absolute magnitude of the dSph.

<sup>e</sup> Heliocentric line-of-sight distance to the dSph in kpc.

<sup>f</sup> Distance from M31 to the dSph in kpc.

<sup>g</sup> Number of successful spectra with  $P_{\text{memb}} > 0.1$ , i.e., dSph member stars.

<sup>h</sup> Systemic velocity of the dSph estimated according to Section 3.3 in km s<sup>-1</sup>.

<sup>i</sup> Total velocity dispersion of the dSph estimated according to Section 3.3 in km s<sup>-1</sup>.

<sup>j</sup> Sources for location, distance, size, and luminosity of the dSph.

<sup>k</sup> Kinematics for these satellites may be moderately affected by contamination due to M31 substructure—see subsection for details.

<sup>l</sup> These measurements depend on an uncertain detection of a cold spike—see subsection for details.

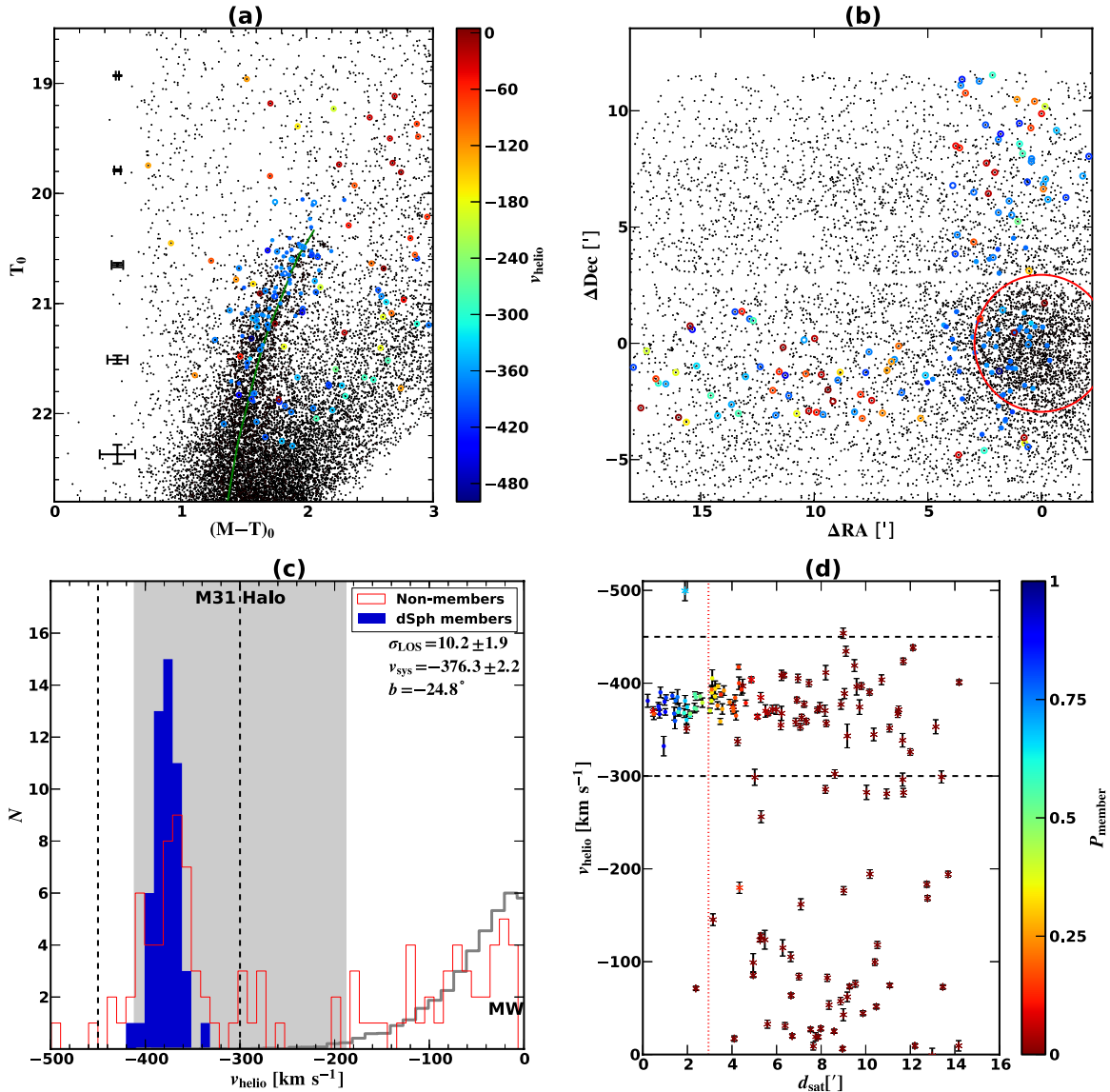
<sup>m</sup> The parameter estimates for this object are based on a different method than the other dSphs—see subsection for details.

**References.** (a) Brasseur et al. 2011; (b) Paturel et al. 2000; (c) McConnachie et al. 2005b; (d) Karachentseva & Karachentsev 1998; (e) Armandroff et al. 1998; (f) Mancone & Sarajedini 2008; (g) Karachentsev et al. 2001; (h) Zucker et al. 2004; (i) Zucker et al. 2007; (j) Martin et al. 2006; (k) Chapman et al. 2007; (l) Majewski et al. 2007; (m) Ibata et al. 2007; (n) Letarte et al. 2009; (o) McConnachie et al. 2008; (p) Martin et al. 2009.

above). This membership approach is thus not sensitive to specific assumptions regarding metallicity or age of the stellar populations and has been found to be successful for both M31 dSphs (Kalirai et al. 2010; Collins et al. 2010) and

MW dSphs (Simon & Geha 2007; Simon et al. 2011). As the plots in Section 4 show, these isochrones closely follow the spectroscopically apparent cold velocity spikes present for most of the satellite fields, validating the approach.

## And I



**Figure 5.** SPLASH results for And I. Panel (a) shows the (Galactic extinction corrected) color-magnitude diagram for And I, and panel (b) shows the spatial distribution of the same objects.  $R_{\text{eff}}$  for the galaxy is indicated in (b) as the large red circle. In both panels, circles are objects with spectra and measured radial velocities. Filled circles are those identified as members according to the criteria discussed in Section 3.2, open circles are non-members, and the color code gives the heliocentric radial velocity. Small points are stellar-like objects detected photometrically, but without spectra. The red points are those that lie inside  $R_{\text{eff}}$  and satisfy the DDO51 pre-selection criteria indicating a high likelihood of being a red giant, while the black points are the remainder. The (green) line passing through these points in panel (a) is the adopted isochrone. Panel (c) shows the heliocentric radial velocities of all spectra from the And I mask with sufficient S/N to measure a velocity as the open (red) histogram. The filled (blue) histogram is for members. Also shown is a gray shaded region indicating the center and  $1\sigma$  width of the M31 halo velocity distributions, as characterized by Chapman et al. (2006). Versions of this figure for other dSphs farther than 100 kpc (projected) do not show this halo distribution, as the typical number of halo stars per mask is very low at those distances. Also shown as a (gray) histogram near  $v_{\text{helio}} = 0$  is an arbitrarily normalized distribution of MW foreground stars derived from the Robin et al. (2003) model in the direction of And I, based on a CMD selection box approximately matched to panel (a). Panel (d) shows  $v_{\text{helio}}$  of all stars with successfully recovered velocities as a function of their distance from the center of the satellite, with error bars derived following the procedure described by Section 3.1. Circles are classified as members and star-shaped symbols are non-members, with the color code signifying the membership probability. The dotted (red) vertical line is  $R_{\text{eff}}$  for And I. The dashed (black) vertical lines in both (c) and (d) indicate the minimum and maximum velocities for inclusion as a member.

(A color version of this figure is available in the online journal.)

The arctan term serves to favor stars that have Na I  $\lambda 8190$  values more consistent with RGB stars than those on the main sequence. This feature is surface gravity sensitive and hence is usually very weak or undetectable in M31 RGB stars, while clearly apparent even on visual inspection of a spectrum of a cool (foreground) dwarf star. This feature becomes weaker for hotter stars, however, and hence only rejects the cooler

MW dwarf stars. The particular functional form is chosen to smoothly transition from 0 to 1 with an adjustable center and sharpness of transition. While the center (here chosen as 2) and sharpness (10) are nominally free parameters, they should not depend on the particulars of any dSph, but rather on the details of the observations and choice of bandpasses for measuring EWs. Hence, we keep them constant throughout this data set.



**Table 4**  
Derived Quantities for M31 dSph

dSph Name <sup>a</sup>	$\log(M_{1/2}/M_{\odot})^b$	$\log(r_{1/2}/\text{kpc})^c$	$\log(L_{1/2}/L_{\odot})^d$
And I	$7.78 \pm 0.18$	$-0.08 \pm 0.02$	$6.35 \pm 0.40$
And III	$7.50 \pm 0.13$	$-0.28 \pm 0.00$	$5.71 \pm 0.12$
And V	$7.53 \pm 0.12$	$-0.35 \pm 0.02$	$5.47 \pm 0.12$
And VII	$8.06 \pm 0.09$	$-0.01 \pm 0.02$	$6.95 \pm 0.12$
And IX	$7.78 \pm 0.20$	$-0.14 \pm 0.02$	$4.87 \pm 0.16$
And X	$6.94 \pm 0.22$	$-0.51 \pm 0.04$	$4.57 \pm 0.03$
And XIII	$6.79 \pm 0.38$	$-0.57 \pm 0.06$	$4.31 \pm 0.16$
And XIV	$7.03 \pm 0.19$	$-0.27 \pm 0.03$	$5.01 \pm 0.03$
And XV	$6.60 \pm 0.36$	$-0.45 \pm 0.05$	$5.55 \pm 0.16$
And XVI	$6.26 \pm 0.74$	$-0.75 \pm 0.04$	$5.31 \pm 0.20$
And XVIII	$7.43 \pm 0.24$	$-0.38 \pm 0.03$	$5.49 \pm 0.03$
And XXI	$7.56 \pm 0.72$	$0.01 \pm 0.03$	$5.36 \pm 0.03$
And XXII	$6.67 \pm 1.08$	$-0.27 \pm 0.03$	$4.11 \pm 0.06$

**Notes.**

<sup>a</sup> Name of the dSph.

<sup>b</sup> Logarithm of the mass within the 3D half-light radius.

<sup>c</sup> Logarithm of the 3D (i.e., deprojected) half-light radius.

<sup>d</sup> Logarithm of half of the dSphs' total luminosity (i.e., luminosity within the 3D half-light radius).

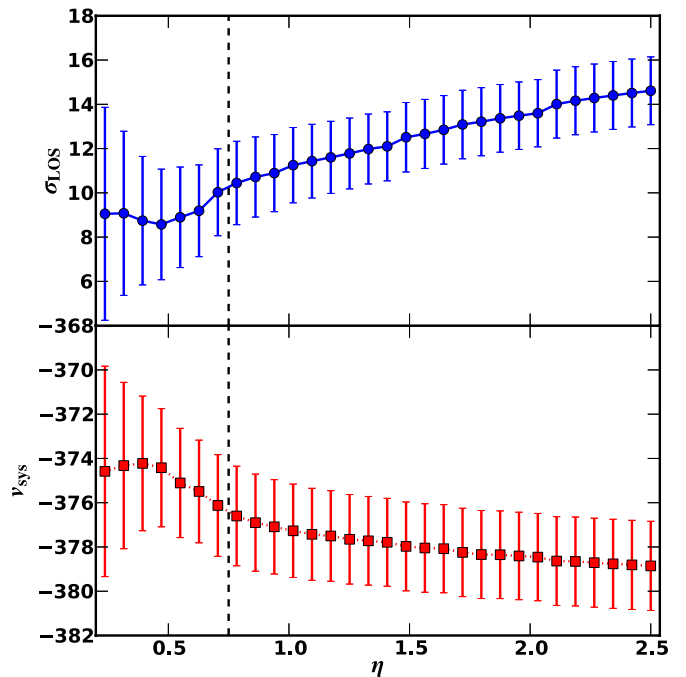
The precise choice of values is informed by Gilbert et al. (2006, Figure 4), as this choice rejects the main locus of MW foreground dwarfs, while only rejecting a small number of giants.

Equation (1) thus defines a membership probability metric that is independent of velocity. Filtering velocity outliers could be warranted because inclusion of a few outlier stars as dSph members that are actually M31 contamination may artificially inflate the velocity dispersion. However, filtering outliers that are true members in the tails of the dSph velocity distribution will serve to incorrectly *decrease* the inferred dispersions.<sup>15</sup> Particularly in the presence of contamination that overlaps all of our observable parameters, it is impossible to completely correct for this effect with the information available. Despite this, we do test applying an iterative  $3\sigma$  clipping filter. We find that it does not affect our membership determination for any dSph other than And I, and we defer the discussion for that particular dSph to Section 4.1. Thus, we adopt a relatively agnostic approach of not using velocities explicitly in our membership formulation, aside from a wide selection window around the dSph's systemic velocity.

The aforementioned membership formulation leads to our adopted definition of “member” stars for the discussion below: those for which  $P_{\text{memb}} > 0.1$ . As is apparent from the velocity histograms in Section 4, this choice is conservative in the sense that it excludes some stars that are likely members based on their radial velocities. However, it also means that those that *are* selected as members are generally rather secure members on the basis of all the information available aside from radial velocity. This serves to decrease the previously noted biases in the kinematical parameters when contamination is small relative to the member population (true in the inner regions for most of the dSphs we examine here).

For cross-checking purposes, we considered alternate selection methods. For And I, III, V, VII, IX, X, and XIV, we considered an alternate member selection method based on velocity,

<sup>15</sup> Numerical simulations assuming Gaussian distributions with sizes like those of our dSph samples imply that this effect can bias the velocity dispersion at the few-percent level.



**Figure 6.** Dependence of kinematical parameter estimates on And I membership parameter  $\eta$ .  $\eta$  is defined as the free parameter of Equation (1) that defines the scale relative to  $R_{\text{eff}}$  outside which membership probabilities falloff. The upper panel (blue) displays the effect of  $\eta$  on the inferred velocity dispersion, while the lower (red) is for the systemic velocity. In both cases the values and associated error bars are derived for each  $\eta$  following the procedure described in Section 3.3. The vertical dashed (black) line is the  $\eta$  value adopted for And I.

(A color version of this figure is available in the online journal.)

spatial proximity, and explicit photometric metallicity estimates (rather than the implicit metallicity dependence based on CMD location described above). We find that the dispersions reported below all agree within  $1\sigma$ . In addition, for a few dSphs, we compare the likelihood method of Gilbert et al. (2006; adjusted to the distance of the dSph instead of the M31 mean). We find that the method used in this paper rejects as MW stars nearly the same set of stars as those labeled probable dwarf stars by the Gilbert et al. (2006) likelihood method.

### 3.3. dSph Kinematical Modeling

With heliocentric radial velocities and membership probabilities, we are now prepared to describe the kinematics of the dSphs of our sample. We model the velocities of member stars in each dSph as a Gaussian distribution with a systematic velocity  $\mu$  and dispersion  $\sigma$ . This assumption of Gaussianity would be violated if there is significant contamination from unresolved binaries (e.g., Minor et al. 2010). However, for the galaxies we present with more than two members, a Shapiro–Wilk test (Shapiro & Wilk 1965) reveals that the null hypothesis that the radial velocity distributions are Gaussian cannot be rejected at the  $p < 0.05$  level for any of the galaxies. This also holds if we apply the test to our entire data set after an offset to the systematic velocity and scaling by the dispersion. Thus, our assumption of Gaussianity is plausible for these data.

We estimate parameters for the radial velocity distribution using a maximum likelihood estimator similar to that of Kalirai et al. (2010) and Walker et al. (2007), but with a factor that includes our method of assigning membership probabilities. The

## And III

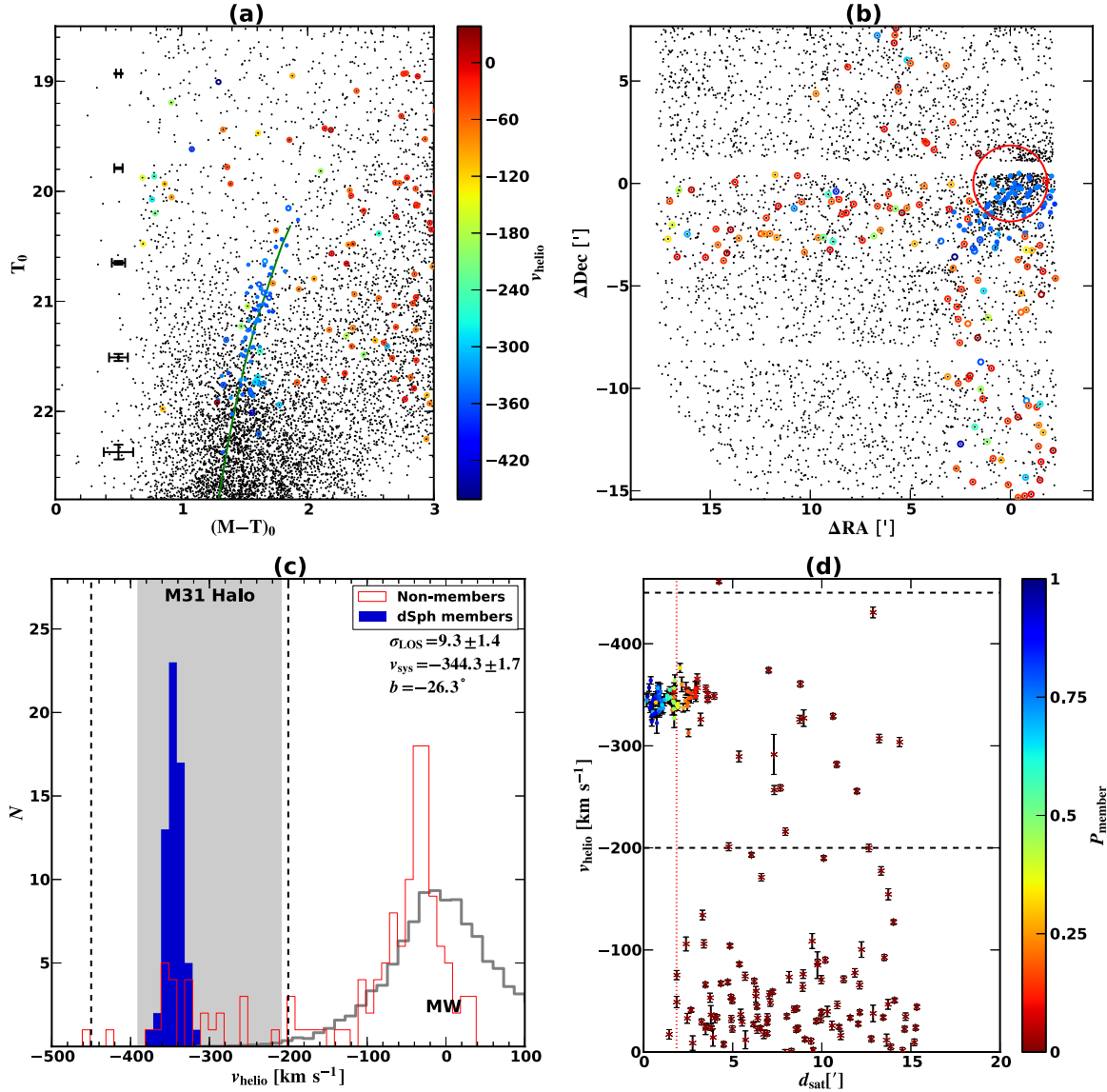


Figure 7. Same as Figure 5, but for And III.

(A color version of this figure is available in the online journal.)

likelihood we adopt has the form

$$\log L(\sigma, \mu | v_i, \sigma_{t,i}, p_i) = -\frac{1}{2} \sum_{i=1}^N \left[ p_i \log(\sigma^2 + \sigma_{t,i}^2) + p_i \frac{(v_i - \mu)^2}{\sigma^2 + \sigma_{t,i}^2} + p_i \log(2\pi) \right], \quad (2)$$

where the  $p_i$  are membership probabilities for each star computed as per Equation (1), and  $\sigma_{t,i}$  is the per-star velocity error (including the  $2.2 \text{ km s}^{-1}$  floor). This likelihood is numerically maximized, and the Hessian is computed at this maximum and inverted to obtain the covariance matrix. We adopt as the error on  $\sigma$  and  $\mu$  the square root of the diagonal elements of this matrix.

As described in the previous subsection, for the final systemic velocity and velocity dispersion, we only accept stars with

membership probabilities  $P_{\text{memb}} > 0.1$ . We thus filter out all of the stars with a high probability of being non-members, because even with low weights from Equation (2), large outliers strongly bias the parameter estimates. For the same reason, we also reject stars with velocities far from the cold spike for each dwarf (limits determined on a case-by-case basis as detailed in Section 4). This serves to filter out a small number of MW foreground stars that happen to lie near a satellite spatially and near the fiducial isochrone in the CMD. The precise range of accepted velocities for each satellite is specified in Section 4. We emphasize, however, that the ranges selected are significantly broader than the velocity peaks for each satellite, so as not to bias the kinematical parameter estimates.

## 4. SPLASH M31 dSphs

With our reduction and analysis procedure outlined in Sections 3.1 and 3.2, we present the full catalog of stars we identify as likely dSph members in Table 2. Non-member stars

are available to the reader on request. We also present a summary of the properties of each dSph studied here in Table 3, including our results for systemic velocities and total velocity dispersions. Dwarfs with no reported  $\sigma_{\text{LOS}}$  have too few likely members to reliably compute a velocity dispersion.

In the subsections that follow, we describe our results and a brief discussion for each satellite in the sample. These each include a figure showing CMDs, spatial distributions, velocity–radius relations, and velocity histograms for each dSph. The subsections also describe any deviation from the fiducial parameters described in the previous section for membership selection. Where relevant, we include a discussion of previous kinematics for each dSph and compare our results.

#### 4.1. And I

And I, originally discovered by van den Bergh (1972), is one of the brightest of M31’s dSph companions ( $M_V = -11.8$ ,  $L_V = 4.5 \times 10^6 L_\odot$ , similar to the MW dSph Leo I). We present our data for this galaxy in Figure 5. A cold spike is immediately apparent near  $-400 \text{ km s}^{-1}$  in the velocity histogram. However, the tails reveal an unusual aspect of And I that complicates the kinematics: it overlaps on the sky with M31’s giant southern stream (GSS; Ibata et al. 2001). The GSS is M31’s largest tidal feature, and while it is  $\sim 100 \text{ kpc}$  behind And I (McConnachie et al. 2003), it overlaps both on the sky and kinematically with And I, as described in Gilbert et al. (2009) and Kalirai et al. (2010). In addition, while the GSS’s overall metallicity and age distribution is distinct from that of And I, the tails of each distribution overlap each other, as well as the general halo population (Gilbert et al. 2009). Thus, the GSS is a potentially major source of contamination, as stars at the boundaries between the distributions cannot be clearly assigned to one group over the other. Additionally, And I is one of the closest dSphs to M31, and the overall morphology of the galaxy hints at tidal disruption (e.g., McConnachie & Irwin 2006). Thus, the choice of the  $\eta$  parameter in Equation (1) has a significant effect on our derived kinematics because GSS stars are likely entering into the sample when  $\eta$  is large.

The primary effect is that increasing  $\eta$  serves to more strongly weight stars at larger radii, which have a higher probability of being GSS contaminant stars. We show how this influences our modeling in Figure 6, which plots how our parameter estimates for  $\sigma_{\text{LOS}}$  and  $v_{\text{sys}}$  are affected by the choice of  $\eta$ . The systemic velocity varies slightly but within the error bars of any given measurement, while  $\sigma_{\text{LOS}}$  varies significantly for large  $\eta$ . We note, however, that the variation is initially relatively flat, suggesting that at small radii the contamination is not affecting the kinematical parameter estimates (at least within the error bars). Hence, we choose a value of  $\eta = 0.75$ , at approximately the value where the upward trend begins. This is significantly more conservative than our fiducial value ( $\eta = 1.5$ ) but serves to down-weight stars in the outskirts of the dSph. We note that while more conservative in the sense of having fewer stars, our result using the analysis procedure presented here is close to (well within error bars of) the earlier analysis of the same data set performed in Kalirai et al. (2010).

Even with a choice of  $\eta$  confining the selection to the core of the dSphs, And I is the only dSph in our sample for which iterative  $3\sigma$  clipping filters any stars. The spectra and CMD locations of these two stars do not show clear indications of being non-members—they show RGB-like spectral features, and they both lie near the fiducial And I RGB in the CMD of Figure 5. Additionally, the velocity distribution is consistent

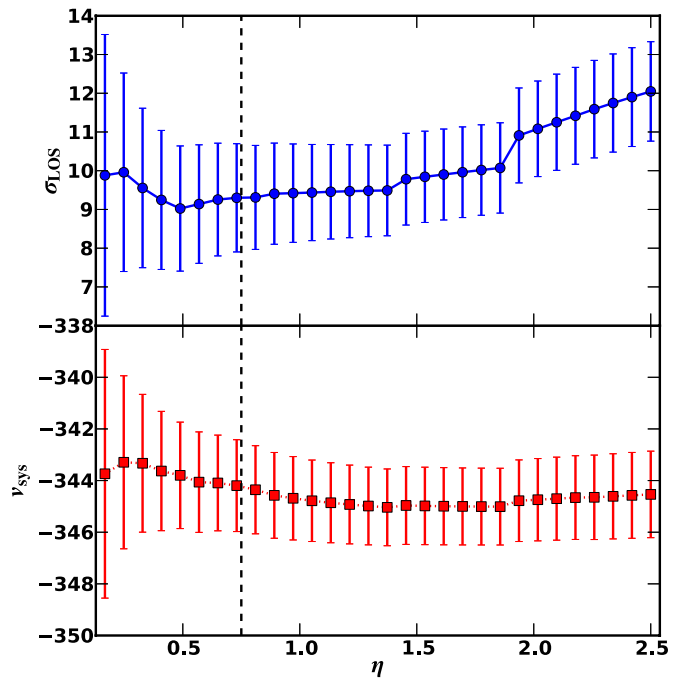


Figure 8. Same as Figure 6, but for And I.

(A color version of this figure is available in the online journal.)

with Gaussian even before sigma clipping,<sup>16</sup> as described further in Section 3.3. Furthermore, the presence of one or two outliers at  $\sim 3\sigma$  is expected for a data set of the size of And I at the 1%–10% level, or for our SPLASH data set as a whole at the  $\sim 45\%$  level. Thus, while it is plausible that these stars are actual members, there is a chance that these stars are M31 contaminants or extra-tidal stars (if And I is tidally disrupting). While we use the unfiltered value in our analyses below for consistency, we also compute the kinematic parameters *without* these stars, finding  $\sigma_{\text{LOS,clip}} = 8.2 \pm 1.7 \text{ km s}^{-1}$  and  $v_{\text{sys,clip}} = -377.0 \pm 1.9 \text{ km s}^{-1}$ . This dispersion is an  $\approx 1\sigma$  offset from the unfiltered value, implying that if the stars are truly non-members, there is a similar level of offset in our mass estimates described below.

#### 4.2. And III

And III is another relatively bright M31 dSph ( $M_V = -10.2$ ,  $L_V = 1.0 \times 10^6 L_\odot$ ), also discovered by van den Bergh (1972). We present our SPLASH results for this satellite in Figure 7. As for And I, a cold spike is clearly present, in this case near  $-350 \text{ km s}^{-1}$ . While near the fringe of the GSS, there is still a hint of possible GSS structure in the form of a slight excess of stars with overlapping velocities far from the center of the satellite. Additionally, there will be a contribution from the tails of the M31 halo distribution given that And III is relatively close to M31 on the sky ( $5^\circ$ , or  $68 \text{ kpc}$  projected). We plot the variation of  $\sigma_{\text{LOS}}$  and  $v_{\text{sys}}$  with  $\eta$  in Figure 8, revealing only a weak dependence, as in And I. However, there is a slight upturn in  $\sigma_{\text{LOS}}$  near our fiducial value. Therefore, we adopt  $\eta = 0.75$ , matching the choice for And I.

Our velocity dispersion for And III here is higher than in Kalirai et al. (2010), who found  $4.7 \pm 1.8$ . This is driven in

<sup>16</sup> Gaussianity was tested for And I using Shapiro–Wilk (Shapiro & Wilk 1965), Anderson–Darling (Anderson & Darling 1952), and Kolmogorov–Smirnov tests, none of which could reject the null hypothesis at the  $p < 10\%$  level.

## And V

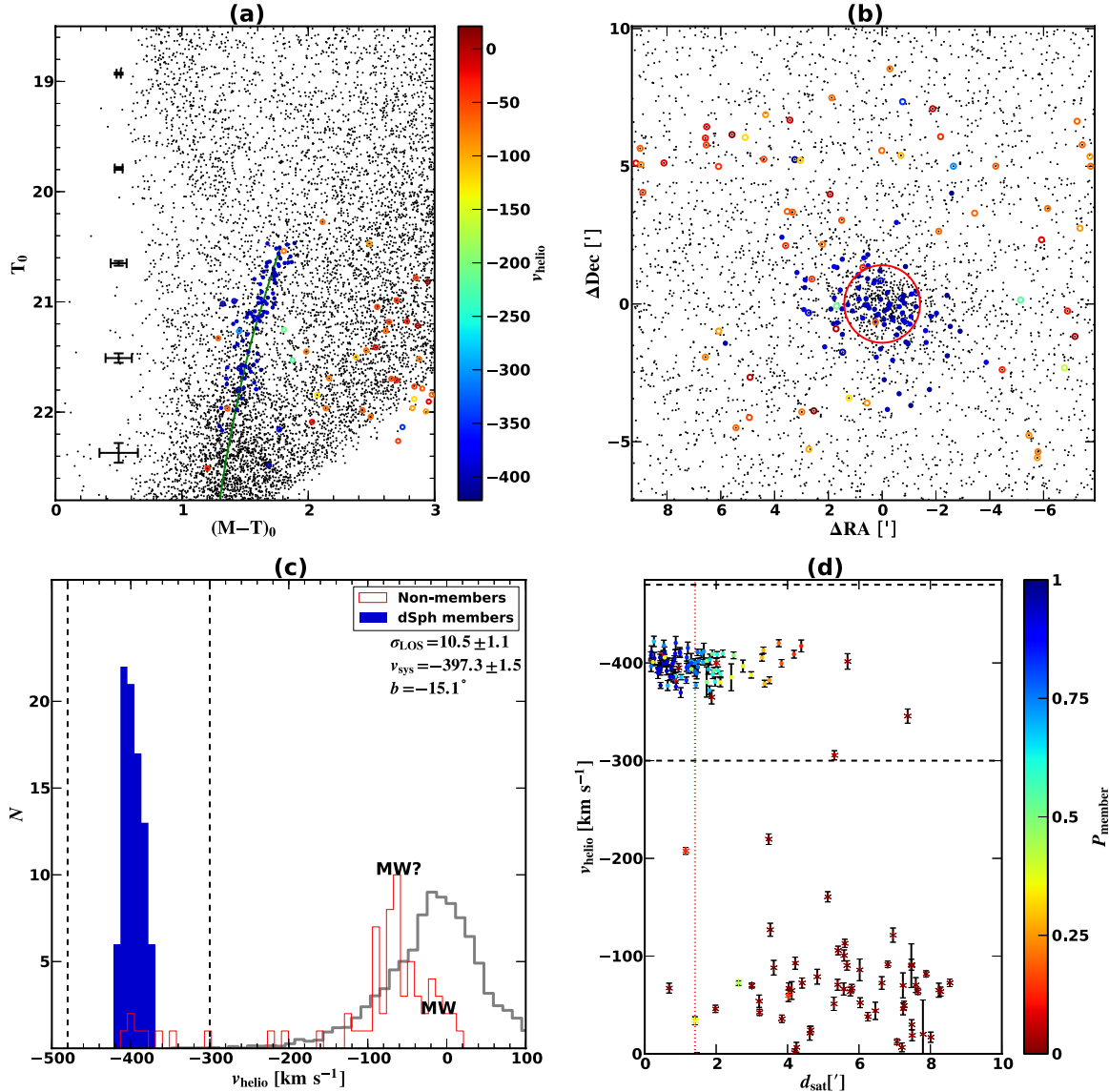


Figure 9. Same as Figure 5, but for And V.

(A color version of this figure is available in the online journal.)

part by the inclusion of a third mask with additional members, and differences in our procedure for determining per-star errors address the remainder of the difference. The primary cause, however, is likely the different methods used for determining membership between this work and Kalirai et al. (2010). In particular, the two most extreme velocity members (the highest and lowest velocity members from Figure 7) are included here and are not in Kalirai et al. (2010). As described in Sections 3.2 and 3.3, such outliers can influence the kinematic parameter estimates, but because of the overlap between the M31 halo/spheroid distribution and the dSph member populations, it is ambiguous whether or not these stars should actually be included as members. For this data set,  $3\sigma$  clipping does not eliminate them, and even with them included, the velocity distribution is consistent with Gaussian (see Section 3.2). Hence, for consistency with the remainder of this paper we keep these stars as members and report the resulting kinematical parameters.

## 4.3. And V

We present our results for And V (discovered by Armandroff et al. 1998) in Figure 9. And V ( $M_V = -9.6$ ,  $L_V = 5.9 \times 10^5 L_\odot$ ) shows a pronounced and very clean cold spike near  $-400 \text{ km s}^{-1}$ , and our kinematic parameter estimates show no significant dependence on  $\eta$ . Because the velocity peak is well away from any contaminants in velocity space, and this dSph is some distance on the sky from M31 ( $8^\circ 0$ , or 110 kpc projected), the membership and parameter estimates are quite secure. Thus, for this dSph, we use our fiducial membership parameters. Additionally, our derived systemic velocity is consistent within error bars with the value adopted in Evans et al. (2000) from the data set of Guhathakurta et al. (2000).

We also note the presence of another possible cold spike in this field at  $\sim -70 \text{ km s}^{-1}$ . At such low heliocentric velocity, this feature is likely to be associated with the MW, particularly given the low galactic latitude of And V ( $b = -15^\circ 1$ ). One possible



## And VII

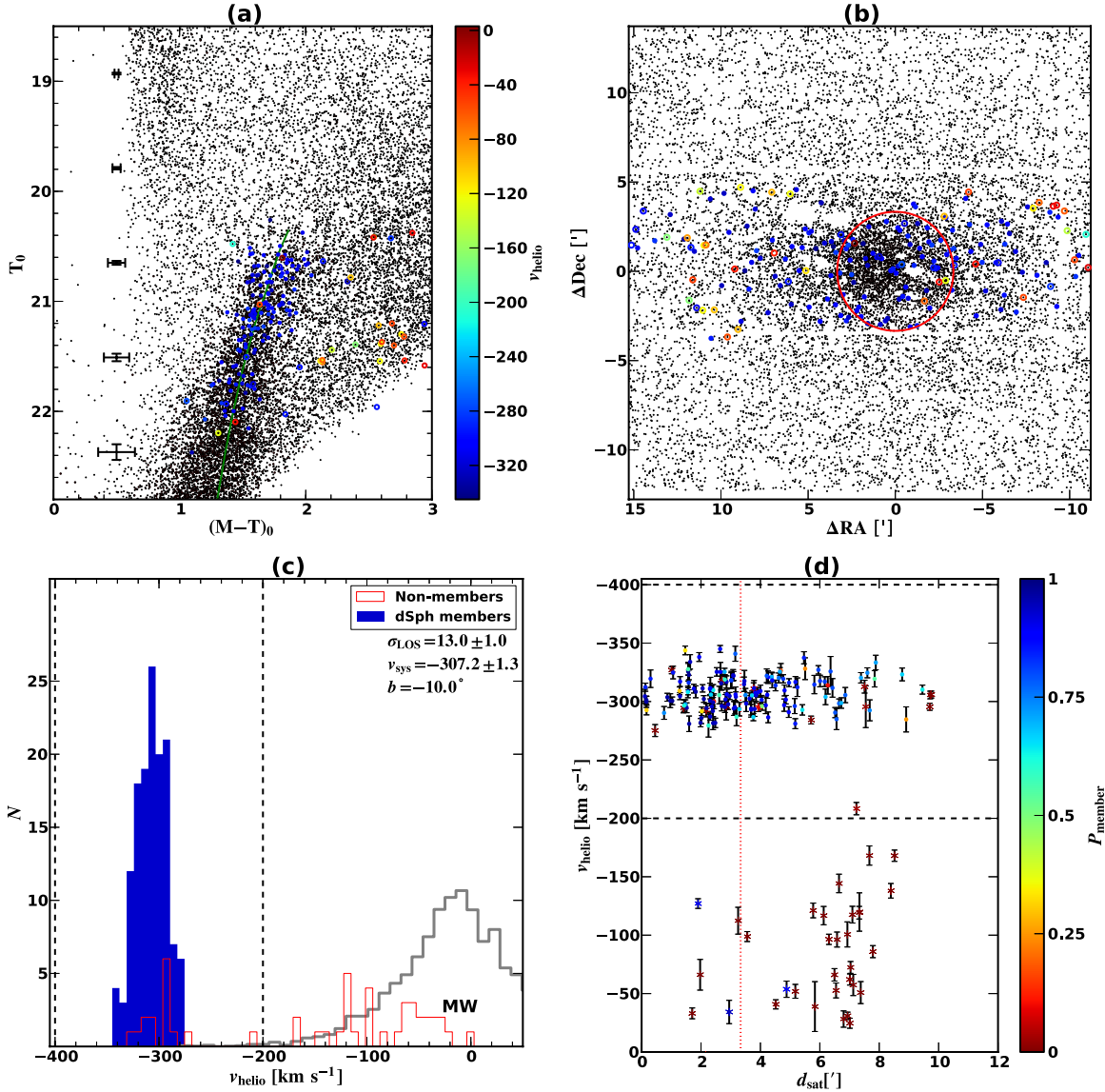


Figure 10. Same as Figure 5, but for And VII.

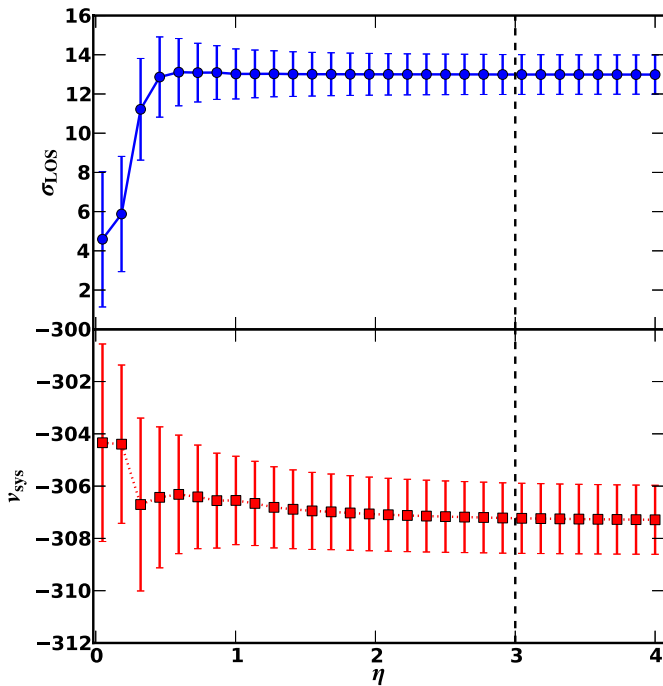
(A color version of this figure is available in the online journal.)

candidate may be the Monoceros ring feature (Newberg et al. 2002), as this overdensity has been noted in fields near M31 (Ibata et al. 2003). The cold feature in Figure 9 has a radial velocity consistent with extrapolation of the velocity–Galactic longitude relation for the Monoceros ring (Crane et al. 2003, Figure 2), but it is not clear that the Mon should have a high density in this field. Alternatively, the cold spike may be related to the Triangulum–Andromeda (TriAnd) feature (described in more detail in Section 4.5), but it is not at the expected radial velocity ( $v_{\text{helio}} \sim -130 \text{ km s}^{-1}$  for this field; Rocha-Pinto et al. 2004). A detailed analysis of the source of this feature is beyond the scope of this paper.

## 4.4. And VII

And VII, discovered by Karachentsev & Karachentseva (1999), is the brightest of M31’s dSph companions, at  $M_V = -13.3$ ,  $L_V = 1.8 \times 10^7 L_\odot$  (although M31’s dE satellites are

only 1–2 mag brighter). We present our results in Figure 10. A cold spike is immediately apparent near  $-300 \text{ km s}^{-1}$ . While this velocity is close to that of the peak of the M31 halo, And VII has a very large projected distance from M31 ( $16^\circ 2$ , or 220 kpc projected), and hence M31 halo contamination is likely negligible for this galaxy. Additionally, as is apparent from Figure 11, our recovered kinematical parameters are nearly independent of  $\eta$ , motivating our choice of  $\eta = 3$ . Because this dSph is so bright, the RGB is very well populated, and the tails of the metallicity distribution are well populated, we increase our isochrone color width  $\sigma_c = 0.25$ . Hence, most of our spectra for these masks are classified as likely members, even well beyond the half-light radius, and these stars cluster tightly in the cold peak. This further underscores the reliability of our DDO51 pre-selection’s ability to select giants over dwarfs when there are enough giants available. Additionally, our recovered  $v_{\text{sys}}$  is almost identical to that of Evans et al. (2000) and Guhathakurta et al. (2000).



**Figure 11.** Same as Figure 6, but for And VII. The  $x$ -axis has been expanded relative to Figure 6.

(A color version of this figure is available in the online journal.)

#### 4.5. And IX

And IX, discovered by Zucker et al. (2004), is one of the fainter M31 dSphs we consider here ( $M_V = -8.1$ ,  $L_V = 1.5 \times 10^5 L_\odot$ ). We present our observations in Figure 12. A cold spike is present near  $-200 \text{ km s}^{-1}$ , but the velocity histogram also shows spikes near  $-330$  and  $-130 \text{ km s}^{-1}$ . The former is consistent with M31 contamination, strong in this field because And IX is quite close to M31 ( $2^\circ.7$ , or 37 kpc projected). This first peak is most likely a mix of smooth halo stars with perhaps a contribution from features such as the NE shelf (Fardal et al. 2007). The spike near  $-130 \text{ km s}^{-1}$ , however, is far from the expected peak for the M31 halo, while the MW distribution for red stars in this field peaks at  $-20 \text{ km s}^{-1}$  (Robin et al. 2003). A possible explanation for this peak is the TriAnd overdensity (Rocha-Pinto et al. 2004). And IX lies near a high-density region of the TriAnd feature, and given the inhomogeneity of the TriAnd feature, a significant TriAnd population in this field is quite plausible. Additionally, this peak is fully consistent with the systemic velocity and width of TriAnd (Rocha-Pinto et al. 2004, Figure 4; And IX is at  $l = 123$  and  $v_{\text{helio}} = -130$  implies  $v_{\text{gsr}} = 46 \text{ km s}^{-1}$  for And IX). An alternative explanation for this feature may be an extension of the M31 disk population, given And IX's proximity to M31. The radial velocity of the peak broadly matches the expectations from Ibata et al. (2005) for this field, but some of the stars in the feature are in parts of the CMD that are not consistent with M31 expectations. Thus, it is likely that this peak is a mixture of TriAnd and M31 disk stars.

Both of the aforementioned peaks are nearly uniformly spread across the field, however, while stars near  $-200 \text{ km s}^{-1}$  are spatially concentrated at the photometric center of the satellite. Thus, we take the peak at  $-200 \text{ km s}^{-1}$  as the peak for And IX, consistent with the results of Chapman et al. (2005). While our membership criteria eliminate most stars in the other two peaks, a small number of outliers nevertheless pass our membership

cuts (while no star list is published in Chapman et al. 2005, this likely includes the outlier mentioned based on the distance given in that work). Thus, for this satellite, we impose more stringent final velocity restrictions (vertical dashed lines in Figure 12) to filter out these outliers, as they are consistent with falling within the distributions of properties for the other populations. We note that kinematic results we obtain with this procedure differ at the  $1\sigma$ – $2\sigma$  level from those of Chapman et al. (2005) and Collins et al. (2010), although our sample has roughly twice the number of member stars.

#### 4.6. And X

Our results for And X (discovered by Zucker et al. 2007;  $M_V = -7.4$ ,  $L_V = 7.5 \times 10^4 L_\odot$ ) are shown in Figure 13. A clear cold spike is apparent near  $-165 \text{ km s}^{-1}$ . This field is significantly farther from the center of M31 relative to And IX ( $5^\circ.6$ , or 77 kpc projected), and hence the M31 halo contamination is much lower. We thus use our fiducial parameters, as changing  $\eta$  does not affect our kinematical results. We note that the dispersion we find here is  $1.4\sigma$  from the result of Kalirai et al. (2009) using the same observations, due to the fact that our method does not formally reject all of the velocity outliers mentioned in Kalirai et al. (2009). If we reject outlier stars in the same fashion, we reproduce a very similar sample as that paper and nearly identical kinematic results.

We also note the presence of another peak in the histogram well offset from the MW expectation at  $\sim -80 \text{ km s}^{-1}$  and a possible smaller peak at  $\sim -130 \text{ km s}^{-1}$ . This field is in a relatively high density portion of the TriAnd feature, and the kinematics of Rocha-Pinto et al. (2004) suggest that the latter peak is due to this feature. However, the broader  $\sim -80 \text{ km s}^{-1}$  peak is well offset from the expected velocity for this feature, suggesting that either TriAnd has a wider velocity dispersion than previously thought or this peak is a different (most likely MW) substructure.

#### 4.7. And XI

Our spectroscopy for And XI (discovered by Martin et al. 2006) is shown in Figure 14. And XI is one of the faintest dSphs in our sample ( $M_V = -6.9$ ,  $L_V = 4.9 \times 10^4 L_\odot$ ), and the field is both small and has few bright RGB candidates. Hence, we do not see an obvious cold spike. We do obtain spectra of two stars near the center of the dwarf that have very close radial velocities. The large distance from M31 ( $7^\circ.5$ , or 102 kpc projected), combined with their proximity to each other and the center of the dSph, renders it plausible that these are dSph members rather than M31 halo stars. Our reported  $v_{\text{sys}}$  is derived assuming that these two are members. However, this result is in conflict with Collins et al. (2010), who report five stars near  $-420 \text{ km s}^{-1}$ . Thus, while we report the mean velocity of these two stars, it is quite possible that these are simple M31 contaminants and we have detected no actual members.

#### 4.8. And XII

Our spectroscopy for And XII (discovered by Martin et al. 2006) is shown in Figure 15. As with And XI, this field is very spare due to the faintness of the satellite ( $M_V = -6.4$ ,  $L_V = 3.1 \times 10^4 L_\odot$ ) and distance from M31 ( $7^\circ.0$ , or 95 kpc projected). We see no clear cold spike, but there are two stars that lie at much more negative velocities than expected for the M31 halo. As with And XI, we report the mean of these two as  $v_{\text{sys}}$ , and again, this is offset from Collins et al. (2010) by

## And IX

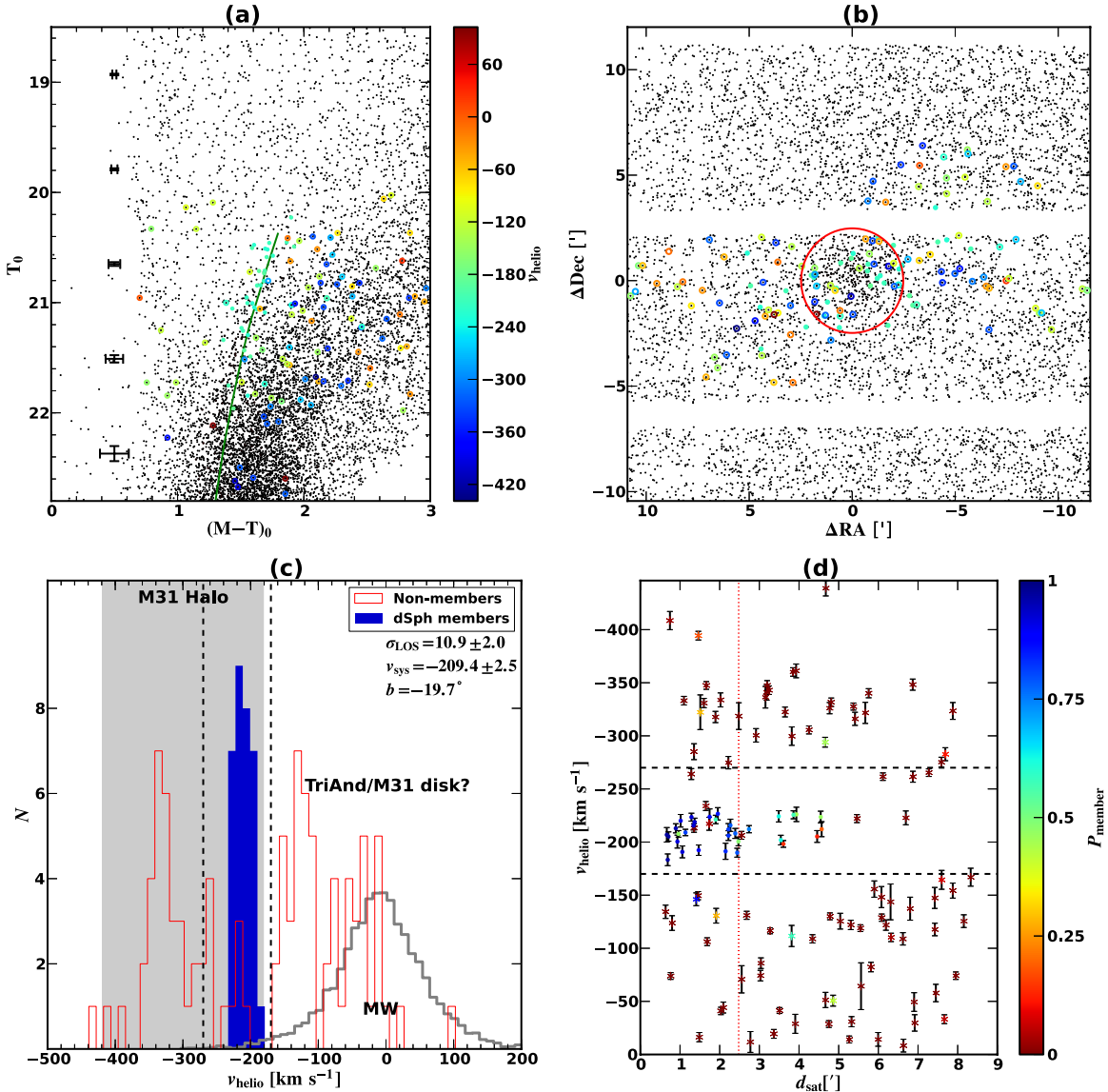


Figure 12. Same as Figure 5, but for And IX.

(A color version of this figure is available in the online journal.)

$\sim 30 \text{ km s}^{-1}$ , so these  $v_{\text{sys}}$  come with the caveat that they may well be non-members.

## 4.9. And XIII

Our spectroscopy for And XIII (discovered by Martin et al. 2006), another faint dSph ( $M_V = -6.7$ ,  $L_V = 4.1 \times 10^4 L_\odot$ ), is shown in Figure 16. We adopt all fiducial parameters for this dwarf and see a cold spike of 12 stars that pass membership cuts near  $-200 \text{ km s}^{-1}$ . The vertical dashed lines in the velocity histogram clearly indicate that we have explicitly cut out stars not near the cold spike, which may at first appear to artificially shrink the dispersion. However, this actually only serves to eliminate two stars (all others failed the membership cuts with the fiducial parameter choices). One of these stars is far removed from the dSph in all three diagrams and hence is likely an M31 contaminant. The other star removed by the cut has  $v_{\text{helio}} = -221 \text{ km s}^{-1}$  and is very close to the And XIII isochrone, as well as lying within the half-light radius. However,

it is also the faintest star in this spectroscopic sample, with an  $S/N = 1.9$ . Hence, only a single CaT line is detected in the spectrum, and it may be biased by nearby sky lines. We therefore reject that particular star and compute the kinematic parameters from the more well-defined cold spike. The resulting parameters are roughly  $1\sigma$  discrepant from the Collins et al. (2010) result, but we note that our sample is four times larger and hence likely more robust.

This field also includes a relatively large number of non-members, despite its large distance from M31 ( $8.5$ , or  $116 \text{ kpc}$  projected) and relatively large distance from the Galactic plane ( $b = -29^\circ$ ). Most of this structure is uniform across the field, suggesting that it is either foreground or large-scale halo structure. In particular, there is a hint of a peak at  $\sim -130 \text{ km s}^{-1}$ , which is consistent with the expected velocities for the TriAnd feature in this field. The stars with radial velocities in the  $-500$  to  $-300 \text{ km s}^{-1}$  range are very unlikely to be of MW origin, however. The distance from M31, assuming the halo model of

## And X

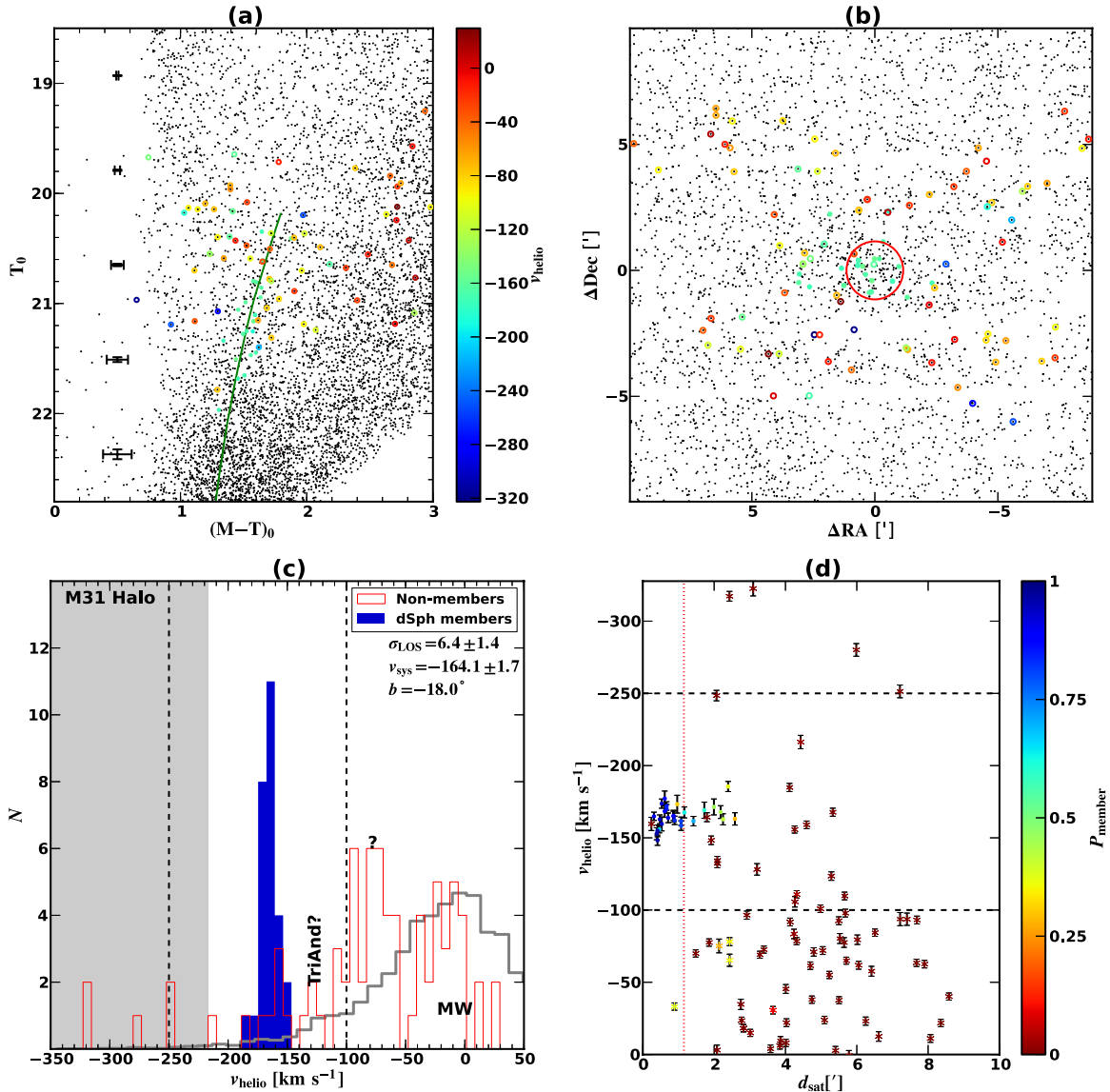


Figure 13. Same as Figure 5, but for And X.

(A color version of this figure is available in the online journal.)

Guhathakurta et al. (2005), implies that the number of M31 RGB stars present in this field is higher than expected, suggesting that the environs of And XIII include some sort of overdensity of M31 stars. Whether this is due to substructure in M31’s halo, tidal stripping of other satellites, or a connection to M33 (And XIII lies roughly along the M31–M33 axis), or simply statistical coincidence, is beyond the scope of this paper.

## 4.10. And XIV

And XIV (discovered by Majewski et al. 2007;  $M_V = -8.5$ ,  $L_V = 2.1 \times 10^5 L_\odot$ ) is unusual because of its large velocity relative to M31. This may imply that it is unbound or barely bound to M31, further supported by its large on-sky distance from M31 ( $11^\circ 7$ , or 160 kpc projected). Because of this isolation, it shows a clean cold peak, immediately apparent in Figure 17 near  $-480 \text{ km s}^{-1}$ . We note that our results here are fully consistent with Majewski et al. (2007), although this should

come as no surprise given that two of our three masks are from the same data set.

## 4.11. And XV

And XV (discovered by Ibata et al. 2007) is a reasonably bright ( $M_V = -9.8$ ,  $L_V = 7.1 \times 10^5 L_\odot$ ), relatively isolated dSph ( $6^\circ 8$ , or 93 kpc projected). Hence, Figure 18 has a readily apparent cold spike near  $-300 \text{ km s}^{-1}$ . However, caution is warranted for this dSph, as it is clear that a significant number of stars in the cold spike do not pass our membership criteria. Examining the spatial diagram demonstrates clearly that there are stars with velocities consistent with the cold spike that are many  $R_{\text{eff}}$  away from the galaxy center. While the velocities of these stars are near that of the M31 halo distribution, some appear to be kinematically nearly as cold as the satellite itself and lie along the dSph locus in the CMD. This suggests that these stars have been tidally stripped from And XV, a conclusion further supported by the dSph’s disturbed morphology (Ibata



## And XI

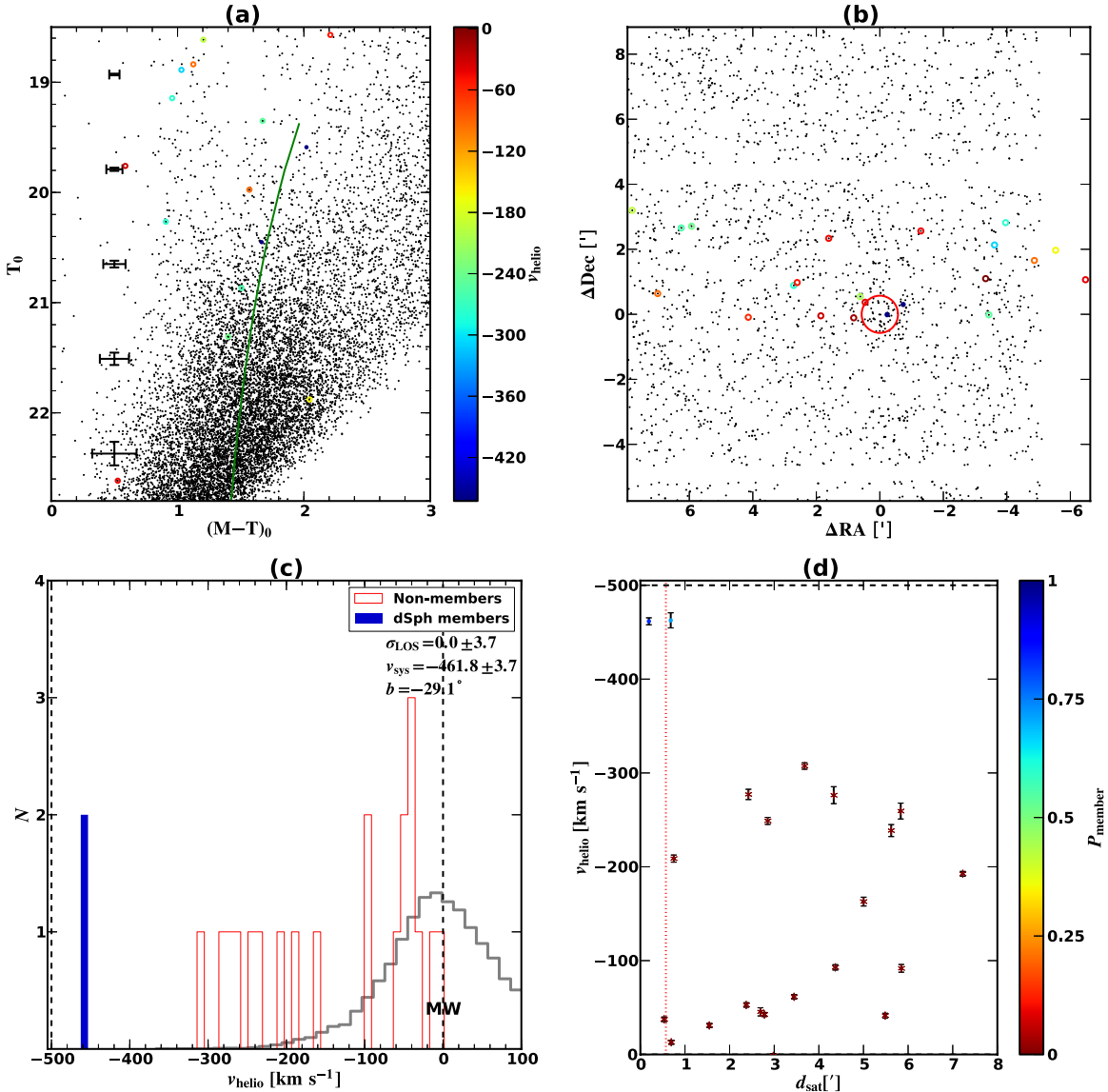


Figure 14. Same as Figure 5, but for And XI.

(A color version of this figure is available in the online journal.)

et al. 2007). Other stars near that velocity are far from the dSph in the CMD, however. These may be stars from the “Stream B” feature of the M31 halo that lies near And XV and has very similar kinematics (Chapman et al. 2008).

Given the presence of these contaminants, it is clear that the choice of  $\eta$  will influence our kinematical results. This is apparent from Figure 19, where the sudden jumps in  $\sigma_{\text{LOS}}$  and increasing trend toward larger  $\eta$  indicate the effect of contaminants. Fortunately, for  $\eta \lesssim 1.3$ , the kinematical parameter estimates are constant (within the error bars). Hence, we use  $\eta = 1.3$  for this dSph. Additionally, we have adopted a value of  $\sigma_c = 0.2$  to account for a somewhat wider CMD for this object (although we note that we obtain similar kinematical parameters if we use our fiducial  $\sigma_c = 0.1$ ). With these parameters, the velocities of the stars we identify as members are stable to iterative  $3\sigma$  clipping and are consistent with a Gaussian distribution (see Section 3.2). Our resulting  $\sigma_{\text{LOS}}$  and  $v_{\text{sys}}$  are  $1\sigma$ – $2\sigma$  discrepant from Letarte et al. (2009), however. This is likely due to a combination of

sample size (our sample is  $2$ – $3\times$  larger), star-by-star errors (ours are smaller due to use of a higher-resolution grating), and our use of a sample that is more strongly weighted toward the center of the dSph. Given the contamination issues outlined above, the stars closest to the center are likely of great importance.

## 4.12. And XVI

Our And XVI (discovered by Ibata et al. 2007;  $M_V = -9.2$ ,  $L_V = 4.1 \times 10^5 L_\odot$ ) field is rather sparse in successful targets, but Figure 20 does show a centrally concentrated cluster of stars near  $-360$  km s $^{-1}$ . Fortunately, And XVI’s relatively large distance from M31 ( $9^\circ.5$ , or 130 kpc projected) and substantial distance from the Galactic plane ( $b = -30^\circ$ ) imply very little contamination. Our fiducial parameters result in a marginally resolved velocity dispersion. The small number of stars (7) implies that this measurement should be taken with caution, however, as excluding a few stars from the analysis typically results in a dispersion formally consistent with zero. We note

## And XII

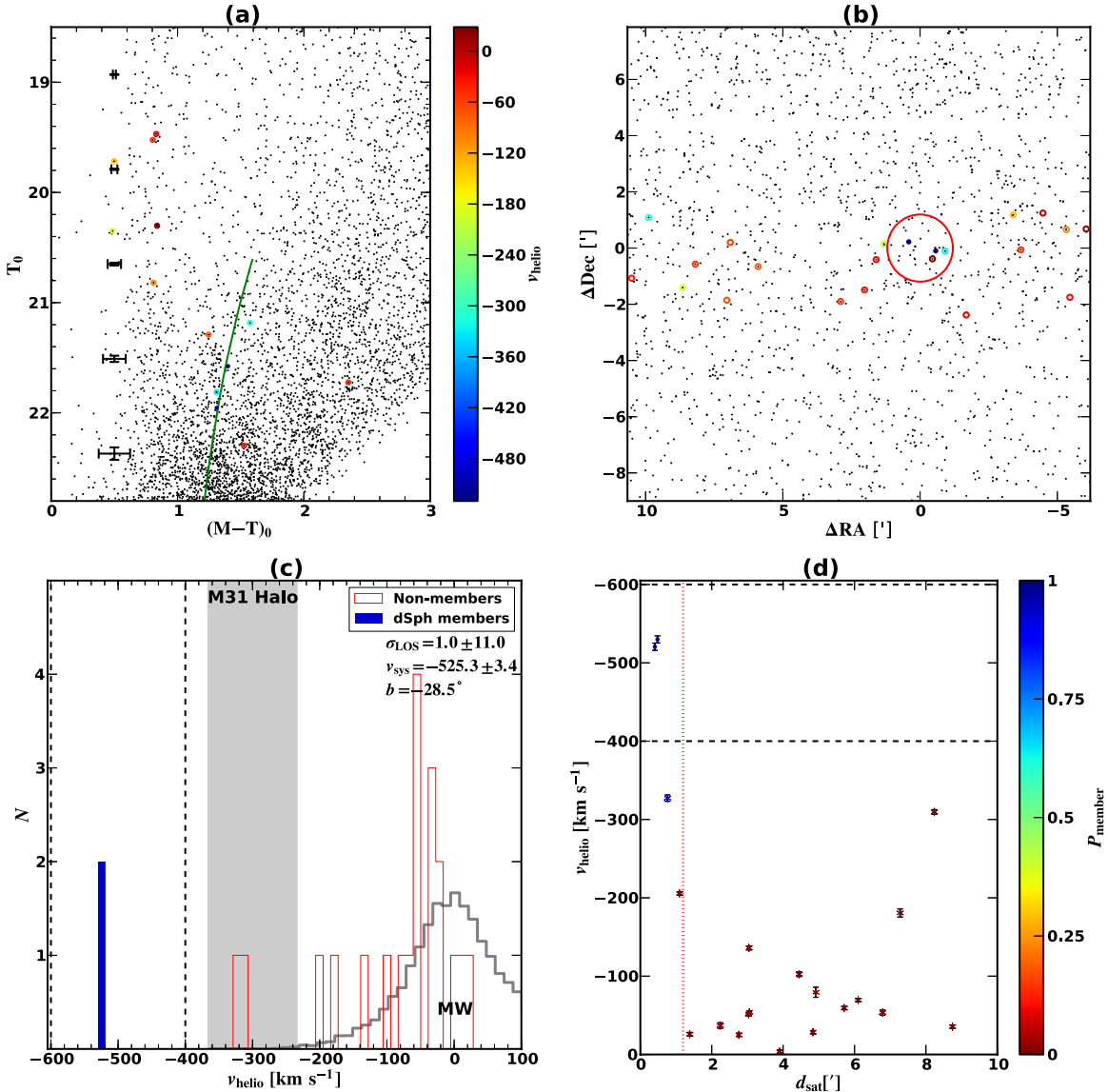


Figure 15. Same as Figure 5, but for And XII.

(A color version of this figure is available in the online journal.)

that our  $v_{\text{sys}}$  is discrepant from Letarte et al. (2009) at the  $\sim 3\sigma$  level, however, even when a dispersion is not resolved.

## 4.13. And XVIII

And XVIII was discovered by McConnachie et al. (2008) and is remarkable in its large line-of-sight distance that formally places it well outside of the M31 system and in fact at the outskirts of the LG. It is also fairly bright ( $M_V = -9.7$ ,  $L_V = 6.3 \times 10^5 L_\odot$ ) and reasonably far from M31 on the sky ( $8.3^\circ$ , or 113 kpc projected), and therefore contamination from the M31 halo will likely be small. We present our results for And XVIII in Figure 21, which shows a clear cold spike near  $-330 \text{ km s}^{-1}$ . These stars are centrally concentrated and far from the typical M31 halo star on the CMD. Hence, we conclude that these observations (the first spectroscopic observations of And XVIII) definitively confirm it as a kinematically cold satellite galaxy. We adopt a  $\sigma_c = 0.3$  to accept the wider-than-typical CMD (primarily due to relatively shallow targeting photometry).

The  $v_{\text{sys}}$  we measure is very close to M31's  $v_{\text{sys}}$ . This is a remarkable coincidence if And XVIII is actually 600 kpc distant from M31 and never interacted. Hence, And XVIII is near its apocenter (and thus nearly at rest relative to M31) and/or closer to M31 (as suggested by our photometry; R. L. Beaton et al. 2012, in preparation).

## 4.14. And XXI

And XXI was recently discovered as part of the PAndAS survey and is relatively luminous ( $M_V = -9.3$ ,  $L_V = 4.6 \times 10^5 L_\odot$ ) but has an atypically large half-light radius/low surface brightness (Martin et al. 2009). It is also quite far from M31 ( $9.0^\circ$ , or 123 kpc projected), and hence contamination from the M31 halo is quite low. We present the first spectroscopic observations of And XXI in Figure 22. The field is sparse, but there is a definite overdensity of stars with radial velocities near  $-360 \text{ km s}^{-1}$ .

For membership criteria, we adopt a larger-than-fiducial  $\sigma_c = 0.3$  to account for the relatively shallow  $T$  exposure. This

## And XIII

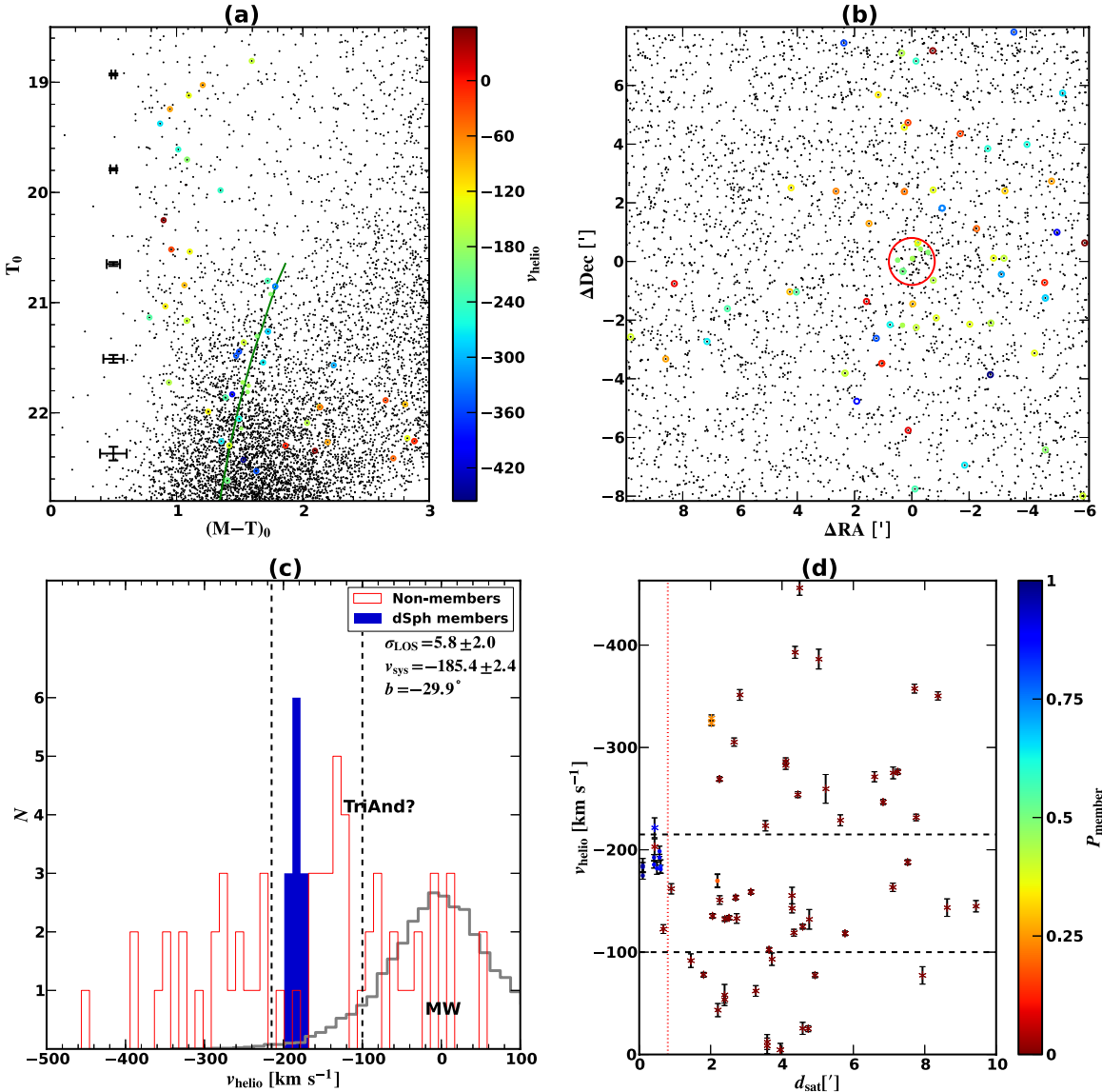


Figure 16. Same as Figure 5, but for And XIII.

(A color version of this figure is available in the online journal.)

mainly serves to include the star near  $M - T \sim 1$  (the faintest in the spectroscopic sample), as that star is very near the center of the dSph and hence is plausibly a member. We also note that the brightest star in the member sample is near (and possibly beyond) the expected tip of the red giant branch (TRGB) for this satellite (Martin et al. 2009). It is also the most distant from the center of the dSph in our sample, hence rendering its membership questionable. It is formally included following our method here, but removing it from our sample results in changes to our kinematic parameters that are well within errors.

With this sample, we find a velocity dispersion with large errors but formally inconsistent (at  $1\sigma$ ) with zero. Additionally, given the clear clustering in radial velocity of four of the stars near the identified RGB of Martin et al. (2009), we consider our data to be spectroscopic confirmation that And XXI is a kinematically cold satellite with  $v_{\text{sys}} \approx -363 \text{ km s}^{-1}$ , although our derived velocity dispersion has large error bars that may be underestimated due to the small sample of only six stars.

## 4.15. And XXII

And XXII, also discovered by Martin et al. (2009), is one of the faintest of the M31 dSphs ( $M_V = -6.2$ ,  $L_V = 2.6 \times 10^4 L_\odot$ ). In Figure 23 we present our results for And XXII. While our success rate at selecting members for spectroscopy was lower for this satellite than others due to a lack of DDO51 pre-selection and the faintness of this galaxy, we do find a concentration of stars within  $R_{\text{eff}}$  that are kinematically cold. These stars lie along a well-defined sequence that matches the CMD of Martin et al. (2009), and the kinematically cold sample shows an elliptical distribution consistent with the photometric measurements (and ellipticity was *not* included in the target selection). Thus, we spectroscopically confirm a cold population consistent with the hypothesis that And XXII is a bound dSph.

Unlike the other dSphs in this survey, the small number of identified member stars for this dSph results in a likelihood distribution of our velocity dispersion parameter estimate that

## And XIV

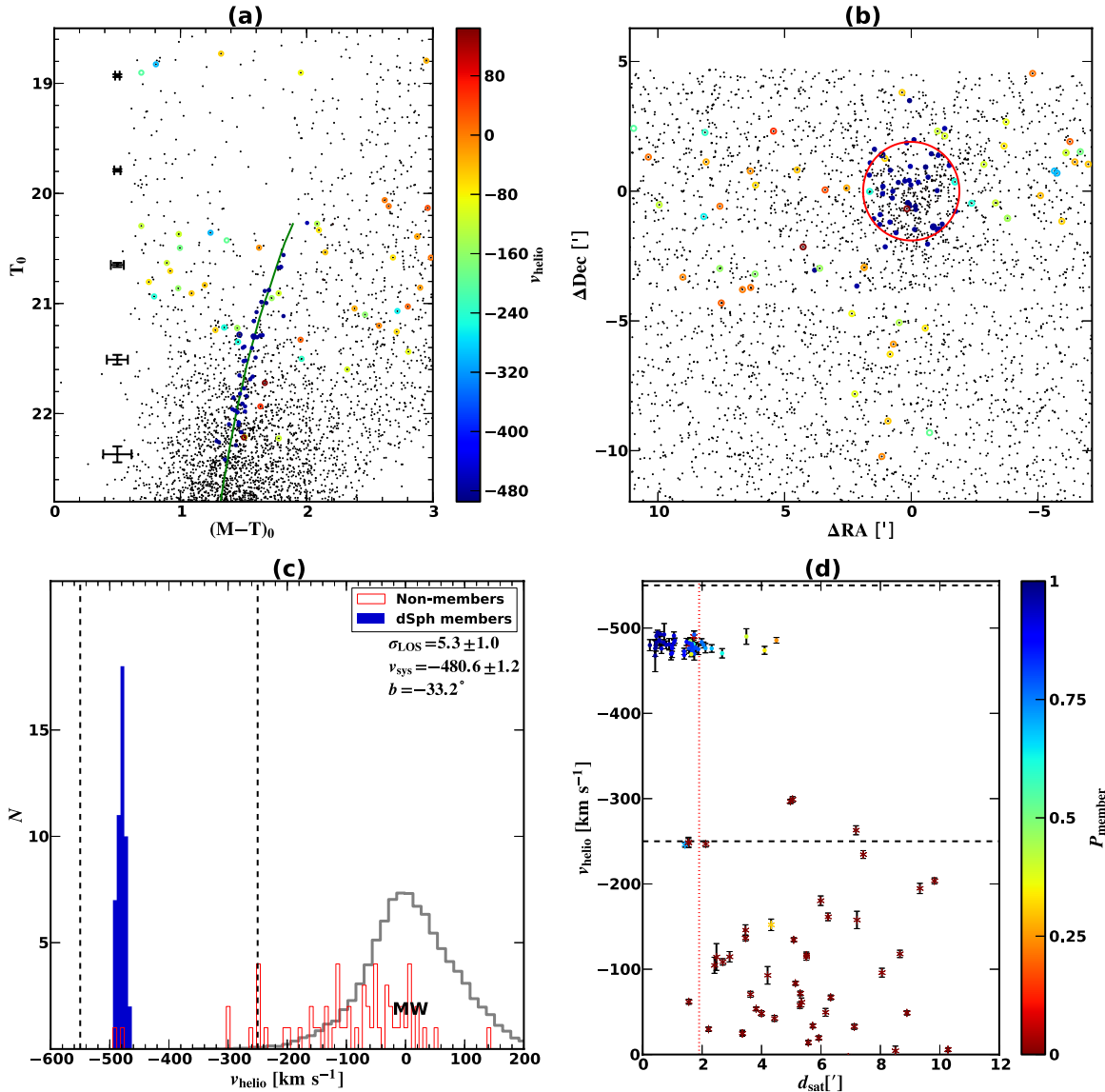


Figure 17. Same as Figure 5, but for And XIV.

(A color version of this figure is available in the online journal.)

is clearly non-Gaussian. Hence, our maximum likelihood technique breaks down because the assumption of normality near the maximum of the likelihood function is invalid. Thus, for this galaxy, we estimate  $\sigma$  and  $v_{\text{sys}}$  using a Markov chain Monte Carlo (MCMC) method with uniform priors on both parameters (Bradford et al. 2011). We have confirmed that the MCMC reproduces the same results as the maximum likelihood method in the case of Gaussian distribution. We quote the MCMC results in Table 3.

This dSph is also much closer in projected distance to M33 than it is to M31: 2:9 and 16:1 to M33 and M31, respectively. This leads to the suggestion that And XXII may be a satellite of M33 rather than M31 (Martin et al. 2009). On the other hand, the line-of-sight distance is closer to M31 than M33, resulting in three-dimensional (3D) distances of 130 and 220 kpc, respectively. Adopting an assumption for the total mass of M31 and M33 based on  $\Lambda$ CDM expectations (Guo et al. 2010, Table 1), the Jacobi tidal radius (Binney & Tremaine 2008) of

M33 (assuming that it is in a circular orbit around M31) is 67 kpc, suggesting that And XXII cannot be bound to M33.

However, our measurements show that And XXII is significantly closer in systemic velocity  $v_{\text{sys}} \approx -127 \text{ km s}^{-1}$  to M33 ( $-180 \text{ km s}^{-1}$ ; e.g., Huchra et al. 1999) than to M31 ( $\sim -300 \text{ km s}^{-1}$ ). Experiments with the Via Lactea 2 (VL2)  $N$ -body simulation (Diemand et al. 2008) suggest that line-of-sight velocities this close for two random subhalos like those M33 and And XXII might inhabit are very unlikely. Furthermore, the line-of-sight distance to And XXII is based on TRGB distance measurements (Martin et al. 2009). The And XXII RGB is sparsely populated, and hence mis-identifying even a few RGB stars near the tip as dSph stars when they are actually M31 halo field stars would bias the distance closer. There is thus a possibility of substantial changes in the TRGB-determined distance with refined stellar membership (e.g., Letarte et al. 2009), in the sense that corrections are likely to push it closer to M33. If this is indeed the case, our finding that it is also closer in



And XV

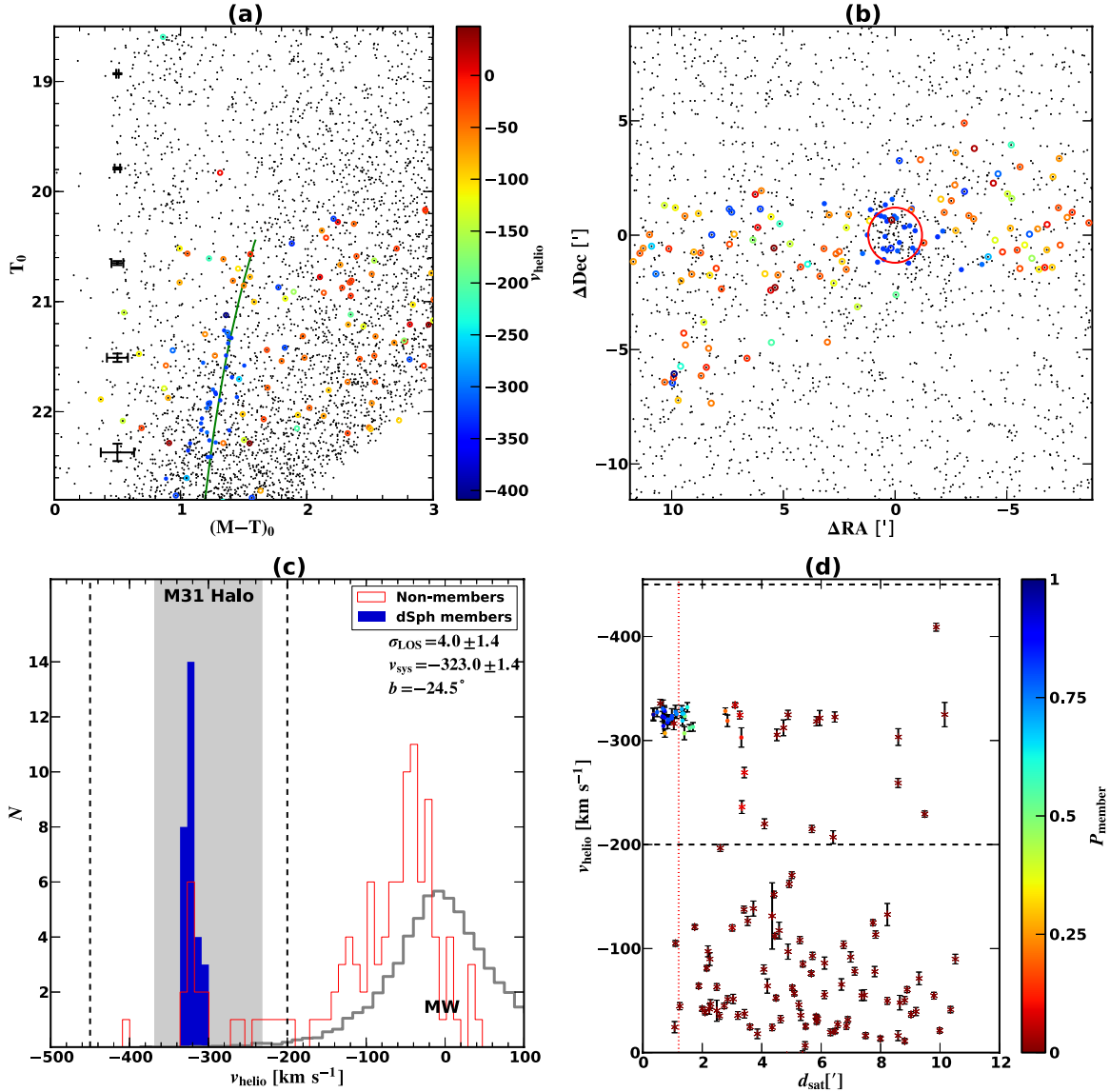


Figure 18. Same as Figure 5, but for And XV.

(A color version of this figure is available in the online journal.)

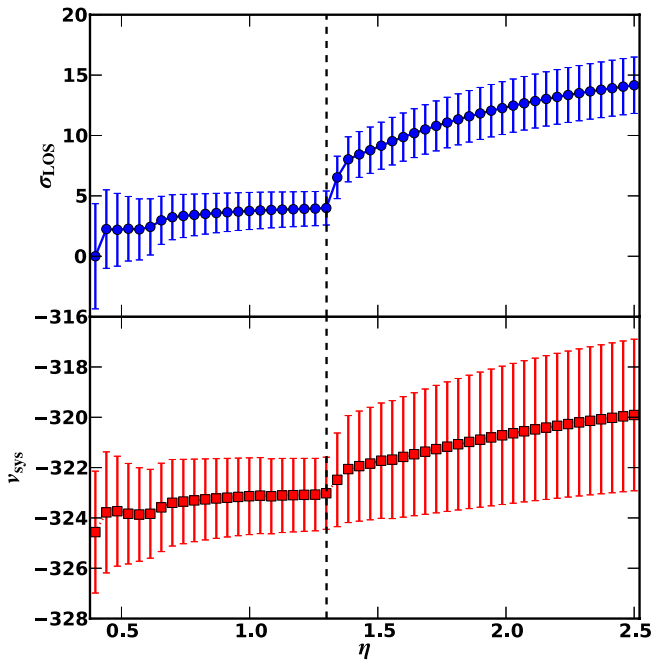
line-of-sight velocity strongly suggests that And XXII is a satellite of M33 rather than M31. In this case, if M33 is taken to be a satellite of M31, as is plausible in a  $\Lambda$ CDM context (Tollerud et al. 2011a), And XXII may be the first detection of a large mass ratio satellite-of-satellite (or sub-subhalo).

## 5. SCALING RELATIONS OF M31 dSphs

With velocity dispersions for most of our sample and structural parameters from a variety of previous studies (see Table 3), we are now in a position to consider where M31 dSphs lie on established galaxy scaling relations. To this end, we consider the scaling relations between mass, luminosity, and size for these satellites. In what follows, we exclude both And XI and And XXI from our sample due to the aforementioned problems estimating their velocity dispersions.

### 5.1. Empirical MRL Scalings

More specifically, we examine the MRL parameter space explored by Tollerud et al. (2011b). This parameter space consists of deprojected (3D) half-light radius ( $r_{1/2}$ ), the mass within that radius ( $M_{1/2}$ ), and the half-luminosity ( $L_{1/2}$ ).  $L_{1/2}$  is straightforward as simply half the observed luminosity of the galaxy.  $M_{1/2}$  and  $r_{1/2}$  are derived from  $\sigma_{\text{LOS}}$  and  $R_{\text{eff}}$  following the prescriptions of Wolf et al. (2010):  $M_{1/2} = 3\sigma_{\text{LOS}}^2 r_{1/2} G^{-1}$ .  $r_{1/2}$  is a simple scaling of  $R_{\text{eff}}$ , at least for light profiles like those of the dSphs we study here:  $r_{1/2} = 1.315 R_{\text{eff}}$ .  $M_{1/2}$  is computed using a standard dynamical mass estimator scaled, crucially, to be the mass within  $r_{1/2}$  and is *insensitive* to the velocity anisotropy. This only holds in the case of a relatively flat velocity dispersion profile, but the dSphs in our data set satisfy that requirement within the error bars (aside from possibly And I and And III—see Sections 4.1 and 4.2).



**Figure 19.** Same as Figure 6, but for And XV. The x-axis has been more densely sampled relative to Figure 6 to better capture the variation near the chosen  $\eta$  value.

(A color version of this figure is available in the online journal.)

These scaling relations thus provide a luminosity and mass, as well as a radius inside which they are applicable, suitable for use with both M31 dSphs from this data set and dSphs from similar MW dSph observations. It is important to note, however, that because  $M_{1/2}$  is derived from the velocity dispersion, any systematic errors in the velocity dispersion resulting from our method of assigning membership to stars will present themselves in these scalings as errors in  $M_{1/2}$ . Furthermore, MW dSphs are likely to have different systematic errors in membership determination due to the quite different contaminant populations. Thus, if these errors are present, they could manifest as systematic shifts in  $M_{1/2}$  for M31 dSphs relative to those of the MW dSphs, even if the populations are intrinsically the same.

With this in mind, we present the three projections of the MRL space in Figure 24. Our M31 dSphs are squares (blue), and we provide the MRL parameters for these satellites in Table 4. We also show the MW dSph compilation of Wolf et al. (2010), with updates for Boötes I and Segue 1 from Koposov et al. (2011) and Simon et al. (2011), respectively, as triangles (red). In the  $L-R$  plane (upper panel), we also plot detection limits for the PAndAS survey (dashed blue; Brasseur et al. 2011) and SDSS (dotted red; Tollerud et al. 2008). Note that the PAndAS detection limits are estimates based on the most marginally detected M31 dSphs, rather than true detection limits from simulations. Given these limits, a fair comparison requires that MW dSphs that could be detected around M31 be compared to the true M31 satellites. Thus, we denote MW satellites within the detection limit by filled triangles and those that could not be detected around M31 as open triangles.

A careful probabilistic M31/MW comparison for the  $L-R$  space (upper panel of Figure 24) has already been performed by Brasseur et al. (2011), and they conclude that there is no significant evidence that the distributions differ between the MW and M31 dSphs. In the lower panels of Figure 24, we see similar levels of scatter as in the upper panel, suggesting there is no

significant deviation. We generate 10,000 Monte Carlo samples of the data in each of the three planes, where each resample is generated by assuming Gaussian distributions for both M31 and the MW data points. We then perform a linear fit, compute the slope and intercept for each resample, and compare the resulting slope/intercept distributions. For both the  $M-R$  and  $M-L$  planes, we find the M31 and MW distributions to be in accord at  $<1\sigma$ , while for  $L-R$  there is a slight disagreement at  $\sim 1.4\sigma$ . Brasseur et al. (2011) find closer agreement with a larger sample and a more sophisticated treatment of errors. In addition, the  $M-R$  relation shows a detectable slope for the M31 dSphs, and thus we perform an error-weighted fit to a power law (dashed blue line in Figure 24), finding  $\log(M_{1/2}) = 2.11 \log(r_{1/2}) + 8.05$ . We conclude that at least our subset of the MW and M31 dSph populations are consistent with lying in the same distribution in major scaling relations.

Despite overall consistency between the scaling relations of the MW and M31 satellites, there are interesting outliers from the overall trends. In the  $M-R$  plane, there are two satellites that show significant deviation from the relation that both MW and M31 satellites seem to follow: And XIV and And XXII. The error bars for And XXII are large due to the poorly constrained velocity dispersion, so its interpretation is unclear. However, And XIV is quite secure and shows a slightly larger  $r_{1/2}$  for its mass than the other satellites (or a small mass for its  $r_{1/2}$ ). And XIV is already unusual in having a large  $v_{\text{sys}}$ , which has motivated the suggestion that it is only now falling into the M31 system for the first time (Majewski et al. 2007). The fact that it is distinct from the other satellites in the scaling relations serves to lend further weight to this idea and supports the notion that environment leaves an observable mark on the structural properties of satellites even in low-density groups like the LG (e.g., Mayer et al. 2001; Kazantzidis et al. 2011; Tollerud et al. 2011a).

In addition, And XV and XVI stand out as outliers from the other satellites in the  $M-L$  plane, in the sense that they are undermassive (for both luminosity and size). For And XVI, the error bars are very large, admitting a reasonable chance that the satellite lies on the relation with the other satellites, but And XV is more secure. An offset for And XV is perhaps not surprising, however, as it shows hints of tidal interaction, as noted in Section 4.11. This is not entirely satisfying, however, as the presence of tidally stripped or halo contaminant stars would typically *increase* the velocity dispersion. While its outlier status may be explained by statistical fluctuation, further investigation of this dSph is needed to resolve this oddity.

## 5.2. Dark Matter Halo Scalings

An additional exercise is suggested by the lower-right panel of Figure 24. The lowest dashed (black) line indicates the line corresponding to a mass-to-light ratio (with  $r_{1/2}$ ) of 1. It is clear that the satellites lie far above this, indicating that they are dark-matter-dominated galaxies with mass-to-light ratios higher than that expected from any reasonable stellar population. This warrants considering what dark matter halos they would be expected to inhabit to give the central densities observed here under the assumption that  $\Lambda$ CDM holds.

We map galaxies onto their dark matter halos by generating a series of Navarro–Frenk–White (NFW) halo profiles (Navarro et al. 1997) and determining the choice of halo that best fits each satellite in the  $M-R$  plane. This approach is described in detail by Tollerud et al. (2011b), and we only summarize here, highlighting the differences. We take the  $M_{1/2}$  and  $r_{1/2}$

## And XVI

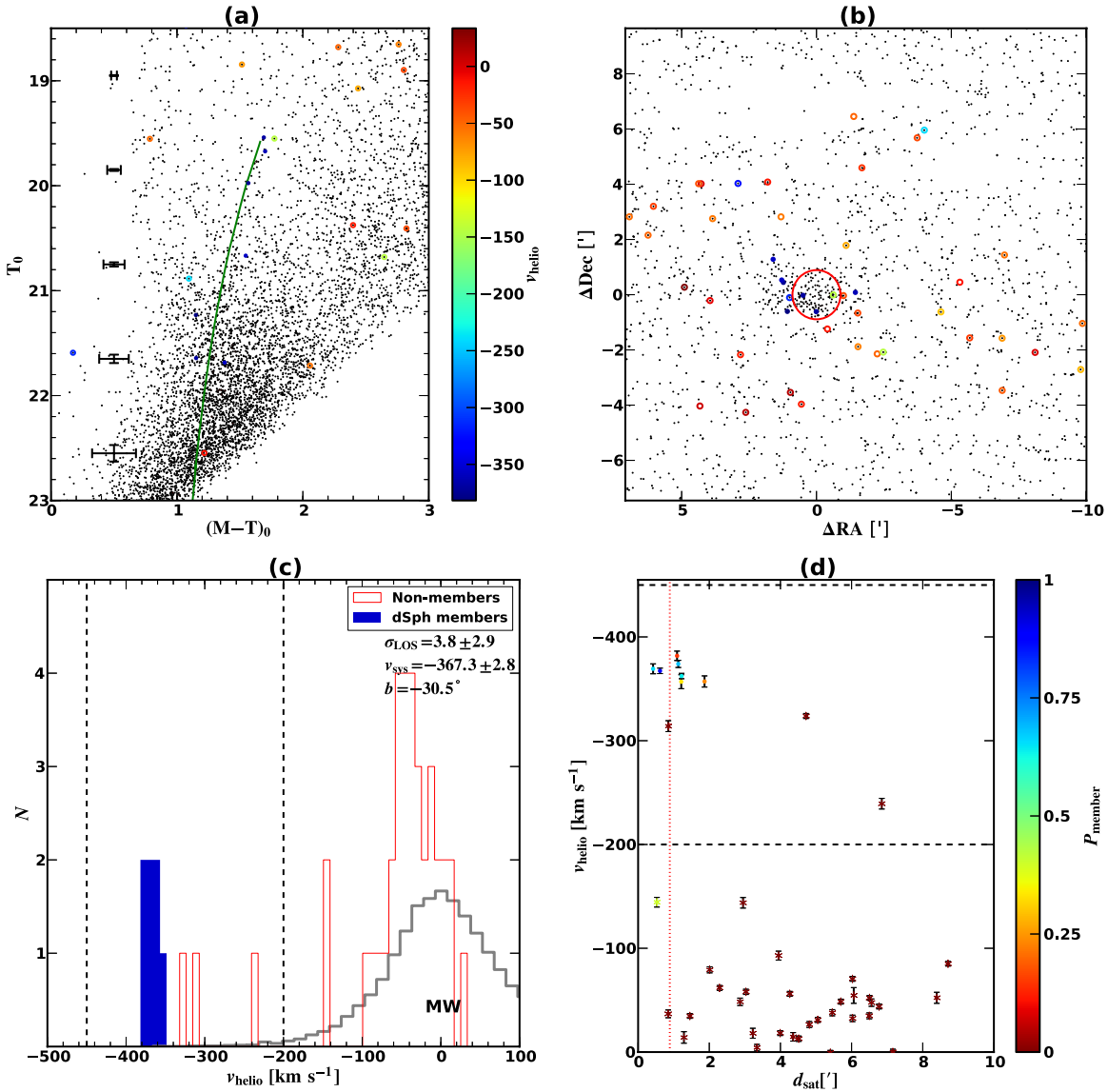


Figure 20. Same as Figure 5, but for And XVI.

(A color version of this figure is available in the online journal.)

estimators from the MRL space to deduce the average density within  $r_{1/2}$  of each dSph studied here. Because  $\Lambda$ CDM dark matter halos are a one-parameter family, we can then map these galaxy scalings onto their dark matter halos by simply matching central densities and reading off the implied dark matter halo's virial mass.

Because we are examining satellites instead of isolated halos, the appropriate dark matter halos to compare with the M31 satellites are *subhalos* of  $M_{\text{vir}} \sim 10^{12} M_{\odot}$  hosts. For subhalos the concept of  $M_{\text{vir}}$  is not always well defined, as their formal virial radii can reach to radii where the host halo is dominant. Additionally, tidal stripping alters the shape and total mass of a subhalo, particularly in its outer reaches. While this stripping does not have as strong of an impact on the central regions where the luminous galaxy sits until the satellite is nearly disrupted (Muñoz et al. 2008; Peñarrubia et al. 2009), it does alter the mass of the dark matter subhalo relative to isolated/field halos. Thus, we use  $V_{\text{max}}$ , the maximum circular velocity of a subhalo (where

$V_{\text{circ}}(r) = GM(<r)/r$ ), as a more stable and well-defined parameter for the mass of a subhalo (e.g., Conroy et al. 2006; Moster et al. 2010; Boylan-Kolchin et al. 2010). Additionally, to reduce NFW to a one-parameter family of models, we use the  $r_{\text{Vmax}}-V_{\text{max}}$  relation from the Aquarius project for *subhalos*, rather than a field concentration-mass relation (Springel et al. 2008; Neto et al. 2007).

It is important to note that this approach depends on the assumption that dissipationless simulations like those of Springel et al. (2008) are sufficient to explain the central densities of satellite galaxies. It is possible that baryons can play a role in altering the densities of dSph halos and this would not be included here. However, as indicated by the lower-right panel of Figure 24, these galaxies have  $M/L \lesssim 10$ . For baryon physics to be a major factor, one would then need to evoke a mechanism of baryonic feedback that displaces significantly more dark matter than the mass formed in stars, and it is unclear what mechanism could achieve this (see, e.g., the discussion in Boylan-Kolchin

## And XVIII

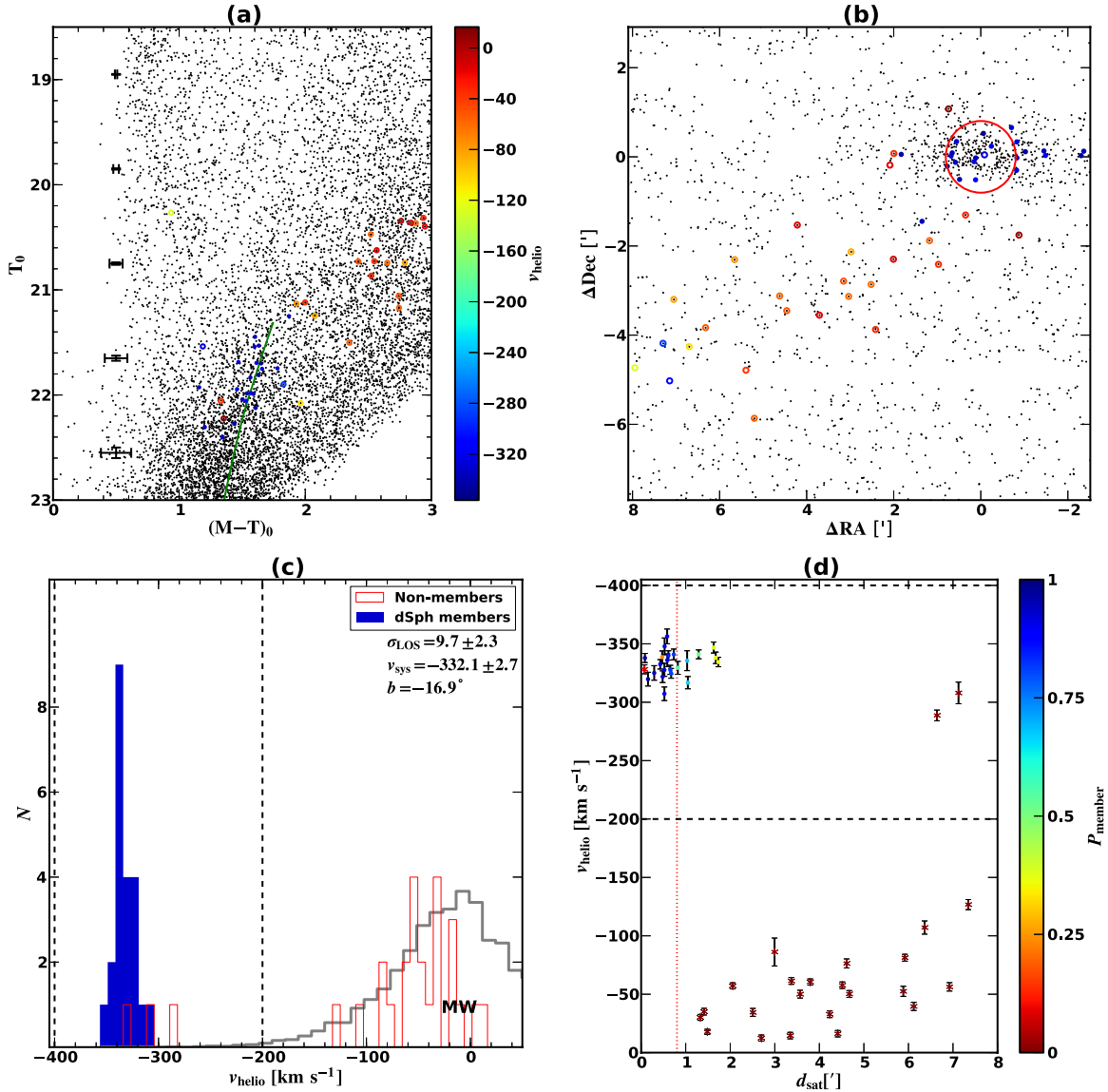


Figure 21. Same as Figure 5, but for And XVIII.

(A color version of this figure is available in the online journal.)

et al. 2012a). Thus, in this discussion we assume that these effects are not significant, as such assumptions have been the backbone of all relevant theory work to date. This allows for direct interpretation of the data in the context of those models.

With this method used to map from the  $R-M$  scalings to dark matter halos, we plot the  $V_{\max}$  values for the M31 and MW dSphs in Figure 25. While there is significantly more scatter here than in the MRL relations, this is primarily because the error bars are relatively large, as the  $V_{\max}$  of a subhalo is quite sensitive to the central density around the scales of these dSphs. We perform the same Monte Carlo simulation method described in Section 5.1, but for the  $V_{\max}$  versus  $L$  relation. These show that the slope and intercepts for the MW dSphs are consistent with the M31 distribution at the  $0.7\sigma$  level. Hence, there is no definitive sign that the MW and M31 have disjoint dark matter halo scatterings. Furthermore, the M31 slope distribution is consistent with zero at the  $0.8\sigma$  level, providing no clear sign that dark matter halo mass scales appreciably with luminosity at these scales. *This*

suggests that the same common mass/halo profile for MW dSphs (Strigari et al. 2008; Walker et al. 2009) holds for M31 dSphs as well.

As in the MRL relations, And XIV, And XV, And XVI, and And XXII are outliers in Figure 25. This reveals a possible interpretation of these results: that these dSphs have anomalously small dark matter halos. In the case of And XV this is not necessarily surprising, as the hints of tidal features suggest that it may be heavily stripped. And XXII and And XVI have large error bars that admit masses well above  $10 \text{ km s}^{-1}$ . For And XIV (and the MW dSph Boötes I), the low mass is a puzzle, because it is below  $10\text{--}15 \text{ km s}^{-1}$ , the scale below which atomic hydrogen cooling becomes inefficient. Hence, if And XVI (and possibly XXII, XVI, or Boötes I) has not been disturbed by interactions, the density implies a dark matter halo that should never have formed galaxies following standard prescriptions of galaxy formation (e.g., Benson et al. 2002; Stringer et al. 2010; Kravtsov 2010). Clearly further investigation is warranted.



## And XXI

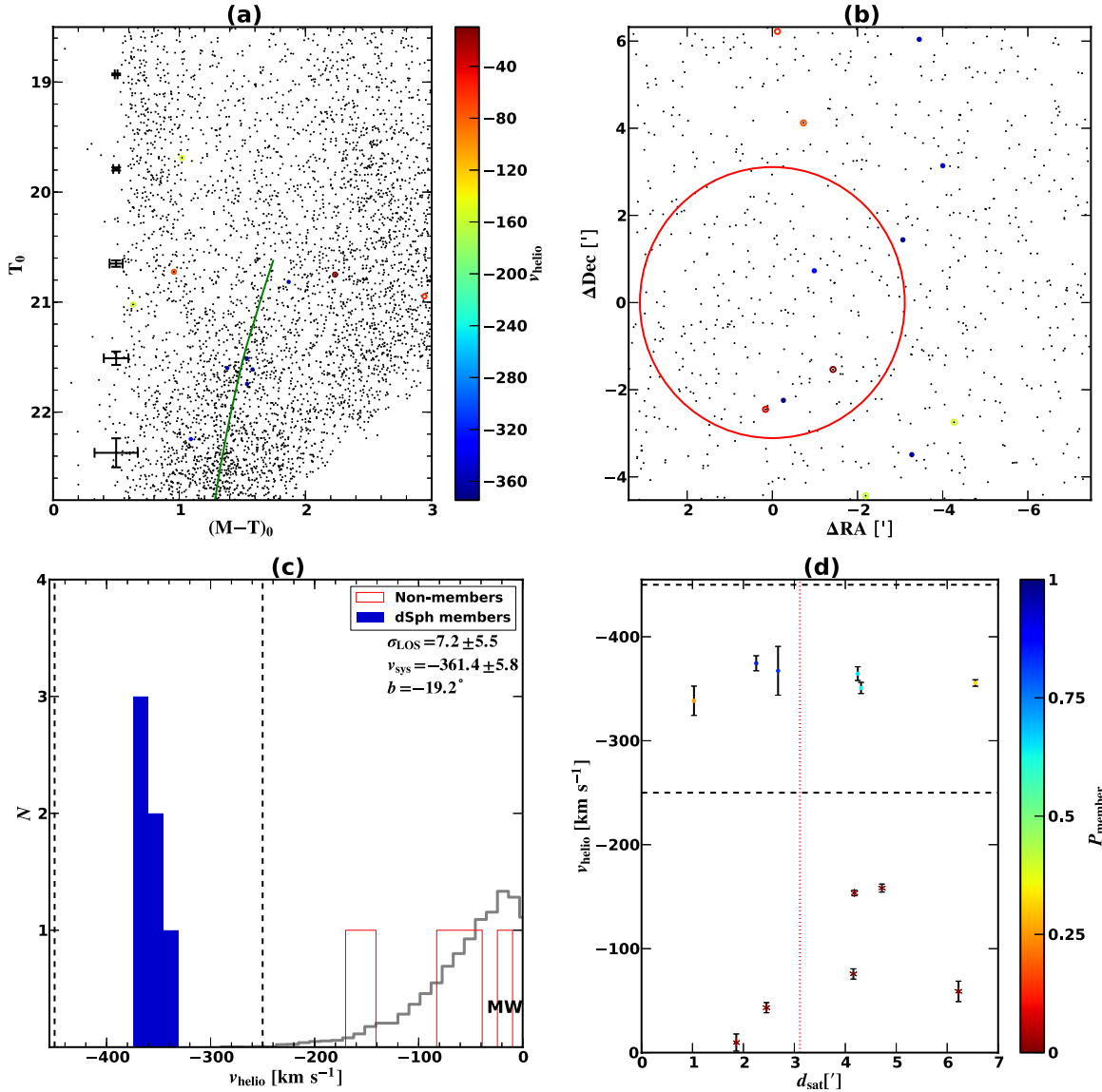


Figure 22. Same as Figure 5, but for And XXI.

(A color version of this figure is available in the online journal.)

## 6. M31 MASS ESTIMATE

In addition to the properties of the dSphs on their own, the kinematics of M31 satellites as a system can be used to estimate the mass of M31 itself (e.g., Evans et al. 2000). While detailed modeling of such systems can provide in-depth dynamical information (e.g., Evans et al. 2000; Klypin et al. 2002; Watkins et al. 2010), here we adopt a simple method, given the relatively small number of tracer particles available here (i.e., the plausible satellites).

We begin by estimating the distribution of  $v_{\text{sys}}$  for the M31 satellites. We use the  $v_{\text{sys}}$  values from Table 3, but with three changes. First, we do not make use of And XVIII given the fact that its distance may imply that it is not a member of the M31 system (see discussion in Section 4.13). Second, we add And II, for which we adopt the  $v_{\text{sys}} = -193.6 \pm 1.0$  from Kalirai et al. (2010) and  $d_{\text{M31}} = 185$  kpc from the distance measurement of McConnachie & Irwin (2006). Third, we also add in M31's

dE population: NGC 205, NGC 185, NGC 147, and M32. We obtain  $v_{\text{sys}}$  and distances for the first three from Geha et al. (2006, 2010), while for M32 we adopt values from Evans et al. (2000). With those changes, we use the maximum likelihood technique described in Section 3.3 (Equation (2)) to estimate the parameters of a Gaussian distribution of the dSph's  $v_{\text{sys}}$ . This yields a velocity dispersion for the M31 satellite system of  $\sigma_{\text{sats}} = 114 \pm 19$  km s<sup>-1</sup>. Furthermore, the resulting mean  $v_{\text{dsphs}} = -298 \pm 26$  km s<sup>-1</sup> of the M31 dSph system *as a whole* is consistent with the mean velocity of M31 and its stellar halo (Guhathakurta et al. 2005, 2006; Chapman et al. 2006; Gilbert et al. 2007).

To infer a mass estimate from this distribution, we adopt an empirical approach appropriate for a  $\Lambda$ CDM context. We adopt a mass estimator proportional to the square of the velocity dispersion, which by dimensional analysis should have the form

$$M_{\text{est}}(<r) = c(r)\sigma^2 r / G, \quad (3)$$

## And XXII

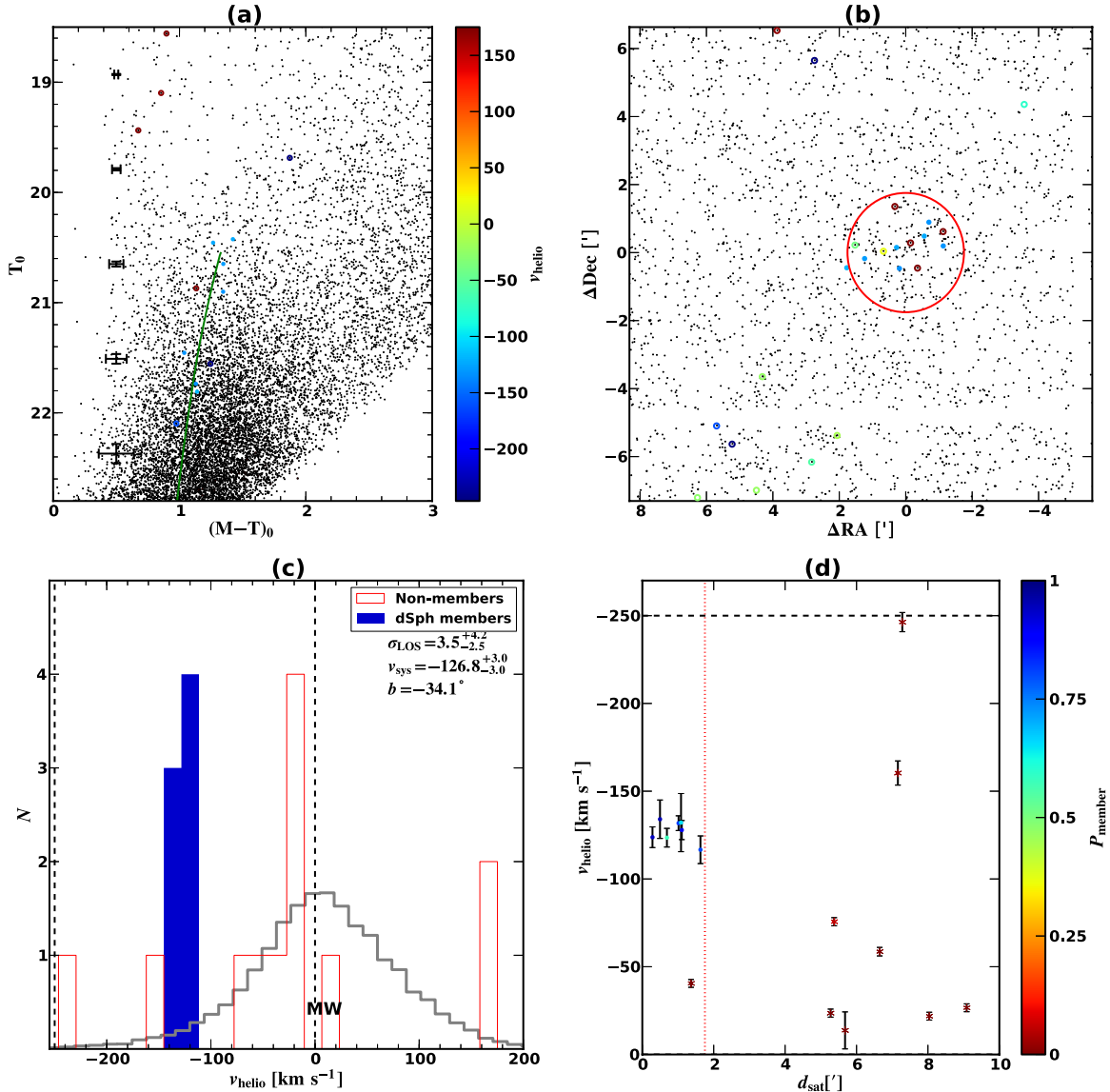


Figure 23. Same as Figure 5, but for And XXII.

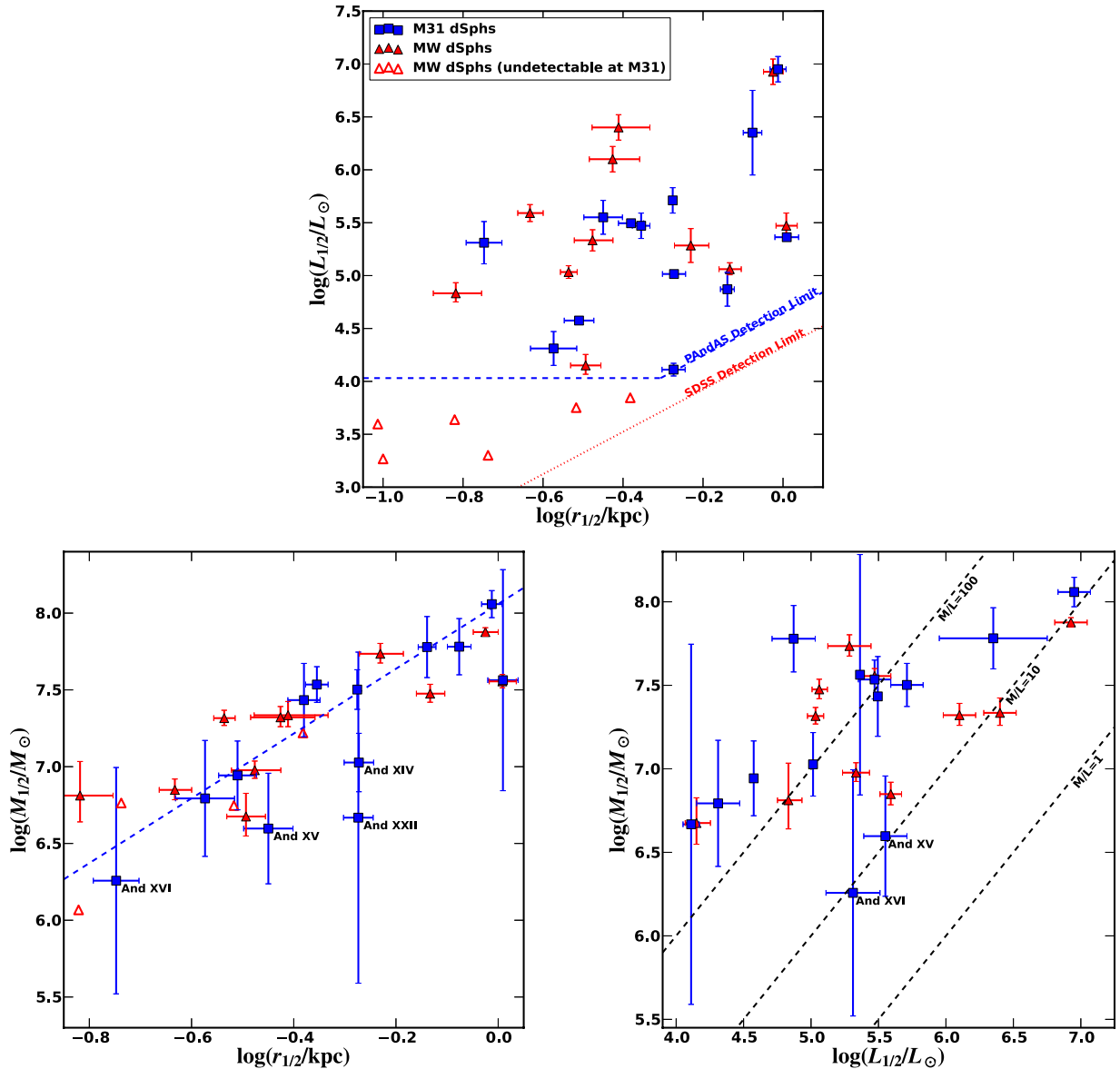
(A color version of this figure is available in the online journal.)

where  $\sigma$  is the velocity dispersion,  $r$  is a radius,  $G$  is the gravitational constant, and  $c$  is a factor to be empirically determined.<sup>17</sup> To determine this factor, we make use of subhalos with  $V_{\text{max}} > 5 \text{ km s}^{-1}$  in the VL2 simulation (Diemand et al. 2008). Kinematics of subhalos of the VL2 halo “observed” from a vantage point as far from the center of the simulated halo as M31 is from the Sun provide a plausible sample of proxies for the M31 satellites. For the radius  $r$  within which we measure the mass, we adopt the median (3D) distance from M31 of the satellite sample, 139 kpc. The median is used here to reduce the effect of small number statistics, but our final virial mass estimate is not very sensitive to the choice of distance, as the correction factor adjusts for different measurement distances. We then determine the correction factor as  $c(r) = M_{\text{true}}(<139 \text{ kpc})/M_{\text{est}}(c = 1)$  for 1000 random

orientations of the VL2 halo. The median and 68% tails of the  $c$  distribution are  $c(r) = 1.96^{+0.37}_{-0.27}$ . We note that this coefficient is somewhat different from the analytically derived Wolf et al. (2010) mass estimator of the same form. This difference is a result of the fact that the Wolf et al. (2010) estimator requires that the full set of tracers be available for computing the half-light radius and velocity dispersion, which is not the case here because the satellite population’s completeness is only well defined to the limits of the PAndAS survey.

With this correction factor in hand, we use Equation (3) with this value of  $c$  and propagate errors for  $c$  and  $\sigma$  to provide a mass estimate for M31 within 139 kpc of  $M_{\text{M31}}(<139 \text{ kpc}) = 8.0^{+4.1}_{-3.7} \times 10^{11} M_{\odot}$ . With a mass within a fixed radius, we are in a position to estimate the virial mass of M31’s dark matter halo. We follow the approach of Tollerud et al. (2011b) and use a grid of NFW halos following the *field*  $c$ – $M_{\text{vir}}$  relation of Klypin et al. (2011), choosing one based on the  $M_{\text{M31}}(r < 139 \text{ kpc})$  determined above. This results in a virial mass for M31 of

<sup>17</sup> This estimator is inaccurately described in some contexts as a “virial estimator”; see, e.g., Merritt (1987, Appendix A).



**Figure 24.** Representation of Local Group dSphs in the MRL space of Tollerud et al. (2011b). The upper panel shows the  $L$ – $R$  plane (half-luminosity vs. deprojected half-light radius). Points (blue) are M31 dSphs, with associated error bars. Also shown are detection limits for SDSS searches of the MW (Tollerud et al. 2008) as the red dotted line, and M31 detection limits from the PAndAS survey as reported in Brasseur et al. (2011). Note that the PAndAS limits are estimates, rather than true limits from simulations. Triangles (red) are MW dSphs as tabulated in Wolf et al. (2010), where solid triangles are within the M31 detection limit, and unfilled triangles are not. The same data set is shown in the  $R$ – $M$  plane (mass within deprojected half-light radius vs. deprojected half-light radius) on lower left. The dashed (blue) line in this panel indicates the best-fit power law with slope 2.1. The lower-right panel shows the  $M$ – $L$  plane (mass within deprojected half-light radius vs. half-luminosity), and the dashed (black) lines in the same panel indicate lines of mass-to-light ratios of 1, 10, and  $100 M_{\odot} L_{\odot}^{-1}$ . (A color version of this figure is available in the online journal.)

$M_{\text{vir}} = 1.2^{+0.9}_{-0.7} \times 10^{12} M_{\odot}$ . This corrected result is comparable to the results of Evans et al. (2000) and Watkins et al. (2010) and thus suggests a moderately lower mass than expected based on assuming abundance matching, i.e., a monotonic halo mass to galaxy luminosity mapping ( $\sim 3 \times 10^{12} M_{\odot}$  based on Guo et al. 2010).

## 7. CONCLUSIONS

In this paper, we have described spectroscopy of M31 dSph satellites as part of the SPLASH Survey. We filter out MW foreground and M31 halo field contamination to identify M31 dSph member stars and use these data to determine  $v_{\text{sys}}$  and  $\sigma_{\text{LOS}}$  for the satellites. Based on these kinematics, we determine for each dSph the implied mass within the half-light radius,

and (under the assumption that these objects are dark matter dominated) we estimate their dark halo properties. This paper can be summarized as follows:

1. We provide a homogenous spectroscopic survey of 15 M31 dSphs and provide radial velocities of resolved stars in these galaxies.
2. We confirm that And XVIII, XXI, and XXII are kinematically cold and hence likely true satellite galaxies.
3. We find that And XXII has a  $v_{\text{sys}}$  close to that of M33, suggesting that it is associated with M33 rather than M31. If so, this is likely the first large mass ratio sub-subhalo (or satellite of a satellite) known.
4. We find that the M31 dSphs obey very similar mass–size–luminosity scalings to those of MW satellites.





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